9th International Workshop
on “Neutrino Telescopes”

Les Houches, Juin 20th, 2001

Global Analysis of Neutrino Oscillation Data

GLF

1. Introduction: neutrino oscillation evidence

2. The “standard” 3ν interpretation

3. 3ν oscillations: atmospheric neutrinos

4. 3ν oscillations: solar neutrinos

5. Conclusions

1. Introduction

Neutrino oscillation evidence

Two kinds of observables:

- total neutrino event rates

$\Longrightarrow$ give information on "averaged" neutrino oscillation probability $< P_{\alpha\beta} >$

- neutrino event spectra (as a function of E, L, L/E, or t)

$\Longrightarrow$ give information on $\dfrac{\partial P_{\alpha\beta}}{\partial x}$ $\quad x = \left| \begin{array}{l} E \\ L \\ L/E \\ t \\ \ldots \end{array} \right.$

crucial to assess oscillations unambiguously !!!

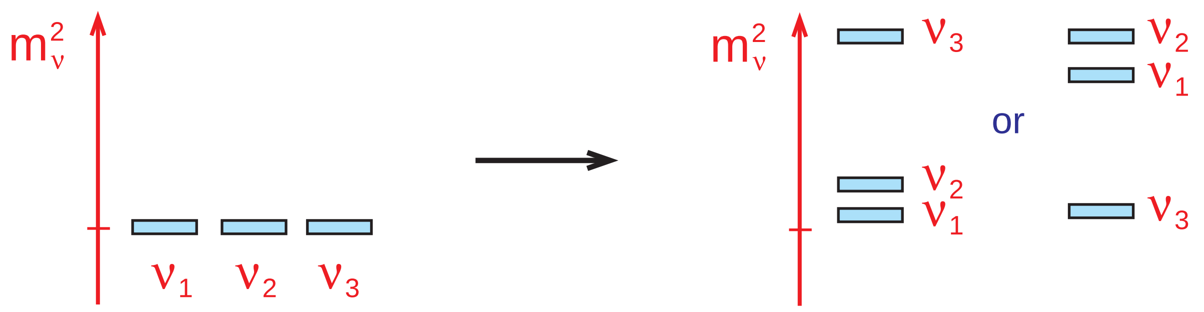
Present evidence

	total rate	spectra
LSND	$P(\nu_\mu \to \nu_e) > 0$ (controversial)	no significant info
solar	$P(\nu_e \to \nu_e) < 1$ (robust)	no significant indication for $\frac{\partial P_{ee}}{\partial(t, E, L)} \neq 0$
atmospheric	$P(\nu_\mu \to \nu_\mu) < 1$ (robust)	$\frac{\partial P_{\mu\mu}}{\partial L} \neq 0$ (very robust) $\frac{\partial P_{\mu\mu}}{\partial E} \neq 0$ (robust)

Limiting our attention to the "robust" information, we can describe it within a "standard" 3ν interpretation

2. The “standard” 3ν interpretation

The two strongest sources of evidence for neutrino oscillations (atmospheric and solar ν anomalies) can be accomodated in a “standard” framework (3ν oscillations) that requires a minimal modification of the electroweak model:



zero mass

at least two nonzero masses

$$U = \begin{vmatrix} 1 & 0 & 0 \\ 0 & 1 & 0 \\ 0 & 0 & 1 \end{vmatrix} \longrightarrow U = U(\theta_{12}, \theta_{13}, \theta_{23}, \delta_{CP})$$

no mixing

at least two nonzero mixing angles (ω, ψ)

We assume then

$$\begin{Vmatrix} \nu_e \\ \nu_\mu \\ \nu_\tau \end{Vmatrix} = U(\theta_{ij}, \delta_{CP}) \begin{Vmatrix} \nu_1 \\ \nu_2 \\ \nu_3 \end{Vmatrix}$$

with in the following

$$\omega = \theta_{12}$$
$$\phi = \theta_{13}$$
$$\psi = \theta_{23}$$

status and prospects of such interpretative framework will be discussed in this talk

Parameters probed by oscillations

(1) $\delta m^2 \stackrel{\text{def}}{=} m_2^2 - m_1^2$

with

(2) $m^2 \stackrel{\text{def}}{=} |m_3^2 - m_{1,2}^2|$

$\delta m^2 \ll m^2$ from phenomenology

$\pm m^2$ both scenarios allowed

(3) $\omega = \theta_{12} \in [0,\pi/2]$

(4) $\varphi = \theta_{13} \in [0,\pi/2]$

(5) $\psi = \theta_{23} \in [0,\pi/2]$

(6) $\delta = \delta_{CP} \in [0,\pi]$

MINIREVIEW

PARAMETER	STATUS	PROSPECTS
m^2	$\sim 3 \times 10^{-3}$ eV^2 within a factor of ~ 2	Good. It will be better and better determined by atm. and LBL expts.
δm^2	$\neq 0$ but multiple ranges allowed (MSW,QV)	Selection of one of the different solutions will take years
ψ	$s_\psi^2 \sim 1/2$ within a factor of ~ 2	as for m^2
ω	$\neq 0$ but multiple ranges allowed (MSW,QV)	as for δm^2
φ	$s_\varphi^2 \lesssim$ few % (CHOOZ) but no reason for $s_\varphi^2 = 0$!	Its determination will be one of the major challenges for future reactor, atmospheric and LBL experiments.
δ	unconstrained. Effects suppressed by $\delta m^2/m^2$	very bad before ν factories

Graphical representation of parameter space

$\delta m^2 << m^2$ implies that:

1. δ_{CP} ~ unobservable — effects doubly suppressed by
$$\begin{cases} \delta m^2/m^2 \to 0 \\ \sin^2\varphi \to 0 \end{cases}$$

we will assume U real. Difficult to prove $\delta_{CP} \neq 0$ in future exps.

2. Solar ν (up to terms of the order $\delta m^2/m^2$ the parameter space is spanned only by three variables:

$$(\delta m^2, \omega, \phi) \Leftrightarrow (\delta m^2, U^2_{e1}, U^2_{e2}, U^2_{e3}) \quad \text{with} \quad U^2_{e1} + U^2_{e2} + U^2_{e3} = 1$$

or equivalently (unitarity)

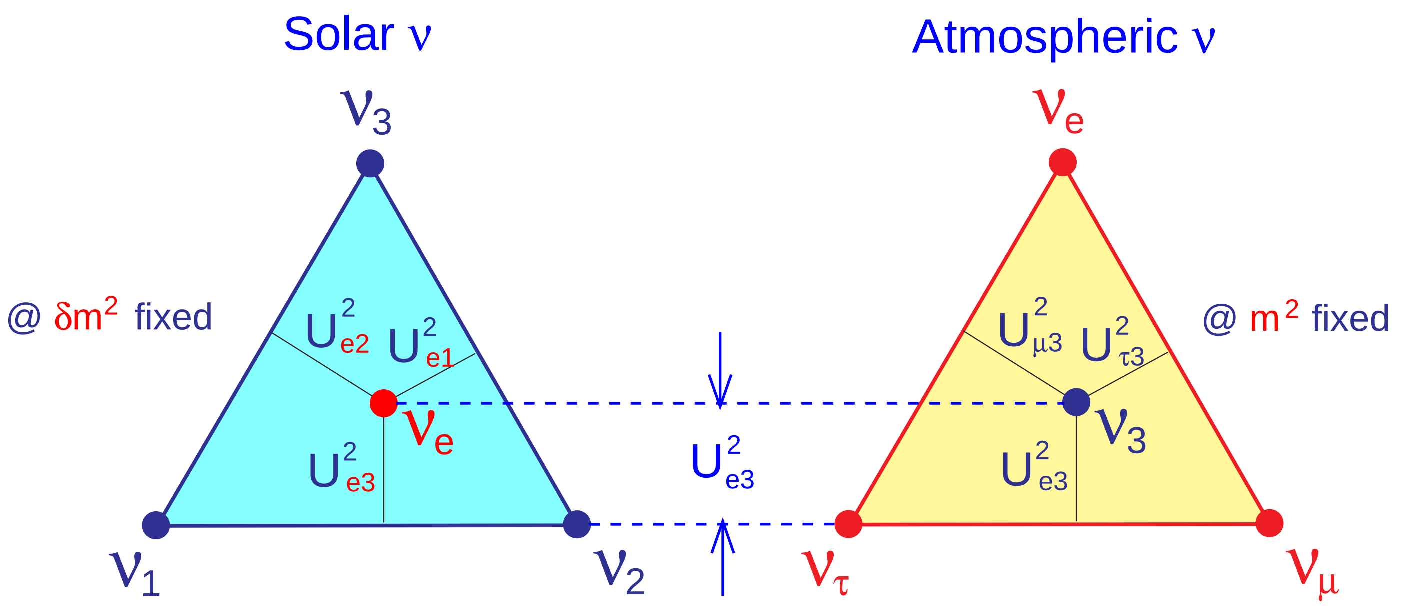
Relevant the mass composition of ν_e:
$$\nu_e = U_{e1}\,\nu_1 + U_{e2}\,\nu_2 + U_{e3}\,\nu_3$$

3. Atmospheric ν (or, in general, terrestrial ν): up to terms of the order $\delta m^2/m^2$ the parameter space is spanned by only three variables:

$$(m^2, \psi, \phi) \Leftrightarrow (m^2, U^2_{e3}, U^2_{\mu3}, U^2_{\tau3}) \quad \text{with} \quad U^2_{e3} + U^2_{\mu3} + U^2_{\tau3} = 1$$

or equivalently (unitarity)

Relevant the flavour composition of ν_3:
$$\nu_3 = U_{e3}\,\nu_e + U_{\mu3}\,\nu_\mu + U_{\tau3}\,\nu_\tau$$



U^2_{e3} probed by solar AND atmospheric ν experiments

A remark:

Beyond the Standard 3ν interpretation

Testing the "standard" 3ν framework for solar+atmospheric ν data does not simply mean refining the measurement of

$$(\delta m^2, m^2, \omega, \psi, \phi)$$

It means also testing and (dis)proving models beyond this framework, which include

- Either new ν states (e.g. ν_s to explain solar+atm+LSND)
- or new ν interactions (FCNC, violation of relativity, ν decay, etc.)

It is particularly important to solve the LSND-KARMEN dilemma in the $\nu_\mu \leftrightarrow \nu_e$ channel, with a new, more sensitive short baseline experiment (MiniBoone)

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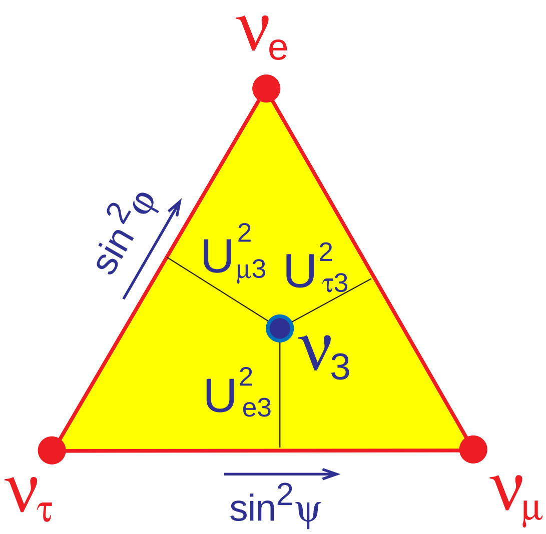
3. 3ν oscillations: atmospheric neutrinos

$$\nu_3 = U_{e3}\,\nu_e + U_{\mu 3}\,\nu_\mu + U_{\tau 3}\,\nu_\tau =$$

$$= s_\phi \nu_e + c_\phi (s_\psi \nu_\mu + c_\psi \nu_\tau)$$

@ m^2 fixed



being

$$U^2_{e3} = s^2_\phi$$

$$U^2_{\mu 3} = c^2_\phi\, s^2_\psi$$

$$U^2_{\tau 3} = c^2_\phi\, c^2_\psi$$

The analysis includes:

- The latest (December 2000: 79.5 kTy) SK data:

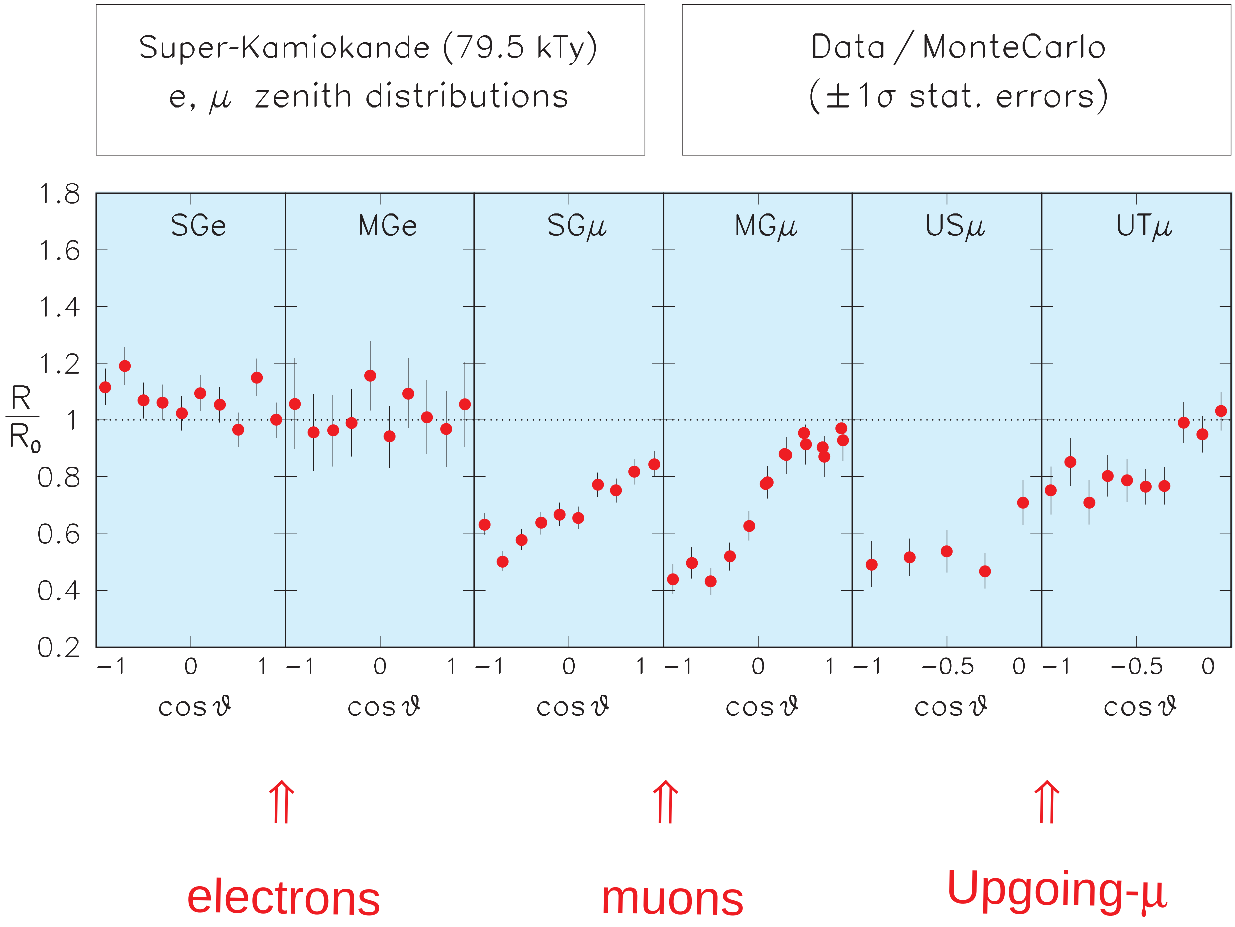
  10 SubGeV e-like bins
  10 SubGeV μ-like bins
  10 MultiGeV e-like bins
  10 MultiGeV μ-like bins
  5 stopping upgoing μ bins
  10 through-going μ bins

- The latest (1999) CHOOZ total rate

  1 data point

SK zenithal distributions (Dec. 2000 data: 79.5 kTy)

(normalized to NO OSCILLATION in each bin)



Comments:

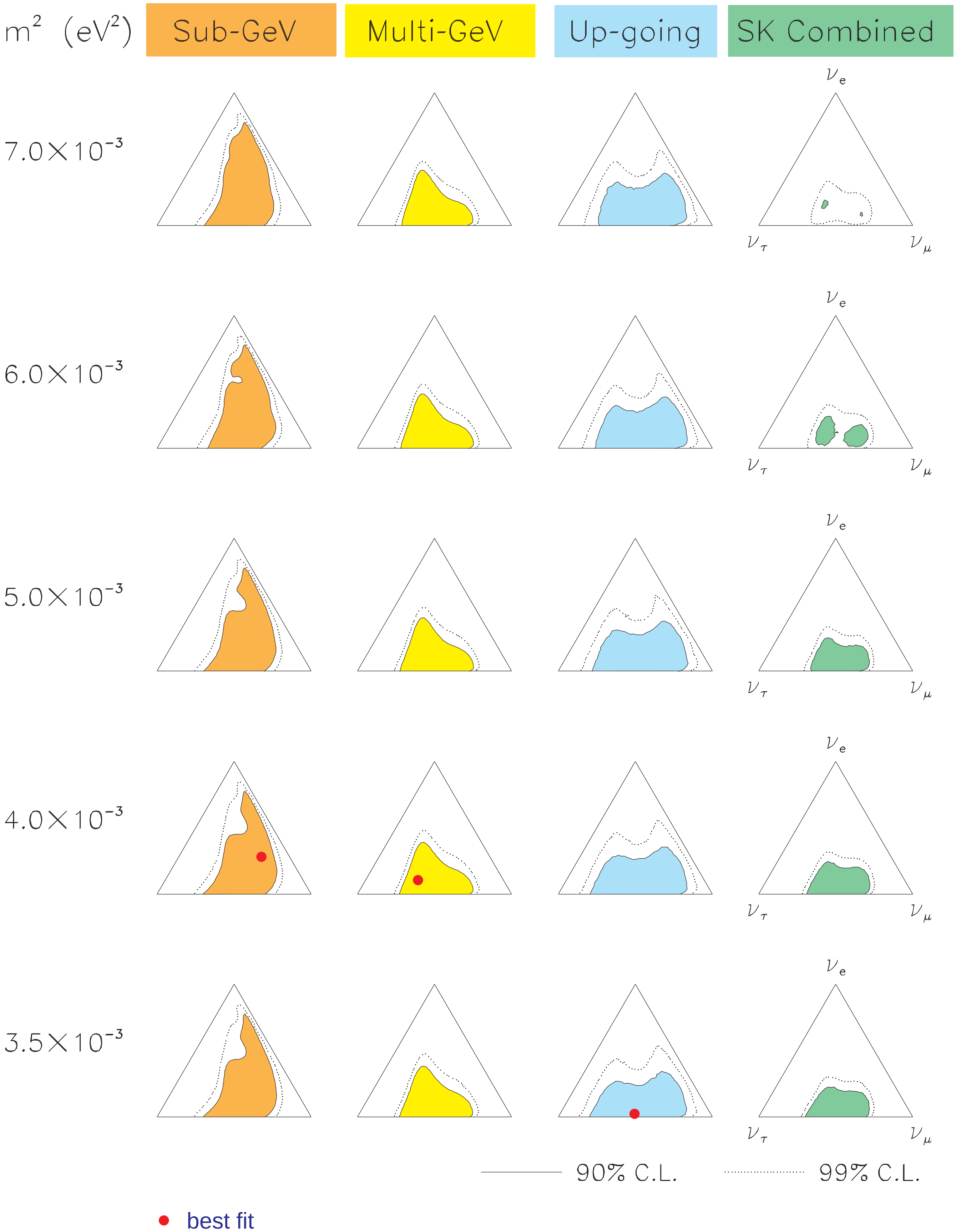
- electrons: no significant deviation from a flat shape (any excess disappeared)
- up-muons: UTμ possibly affected by fluctuations (shape not statistically stable yet)

CHOOZ

$$\frac{R_{\text{exp}}}{R_{\text{theo}}} = 1.01 \pm 2.8\% \text{ (stat)} \pm 2.7\% \text{ (syst)}$$

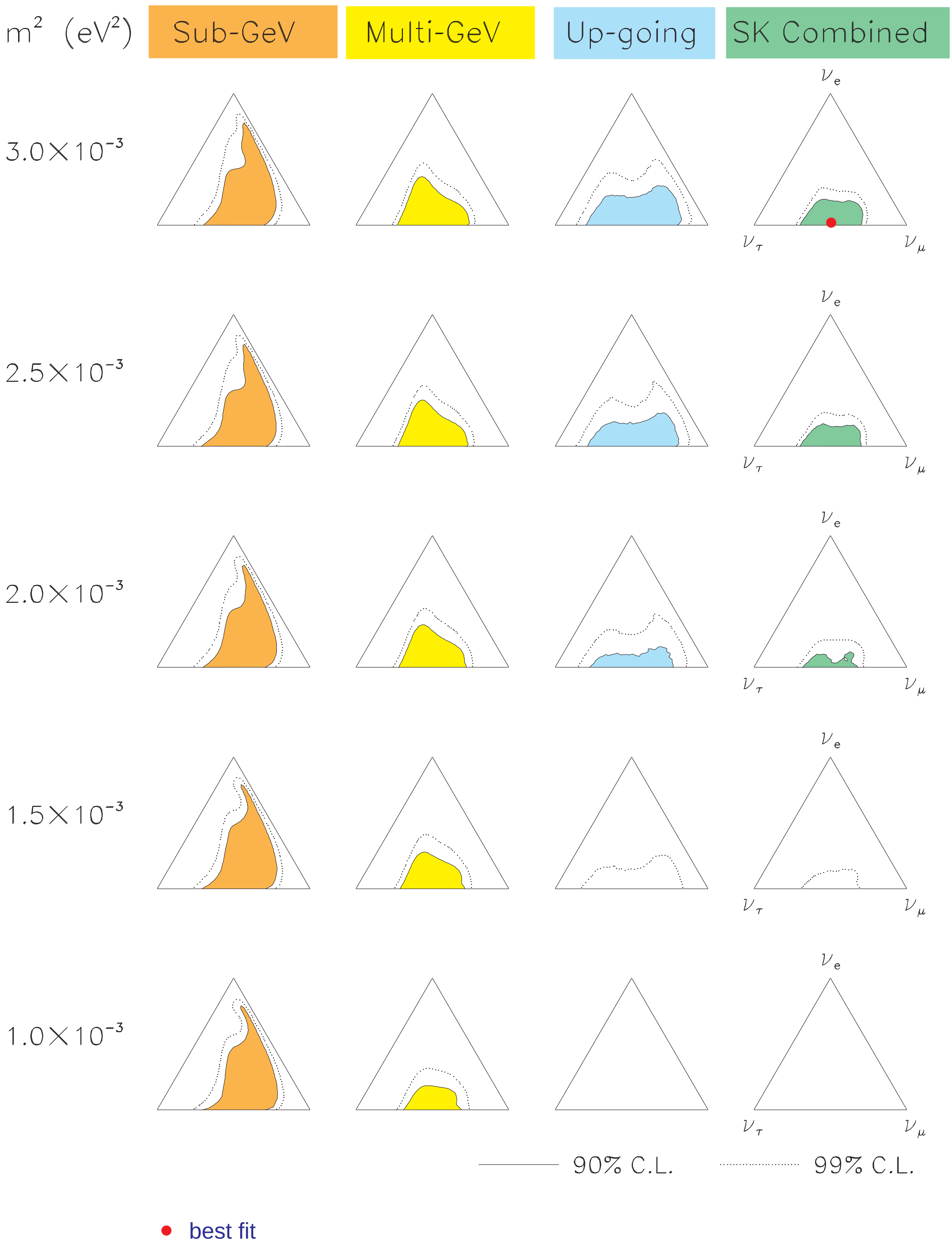
Allowed regions in a three-flavour approach

(Dec. 2000 SK data: 79.5 kTy)



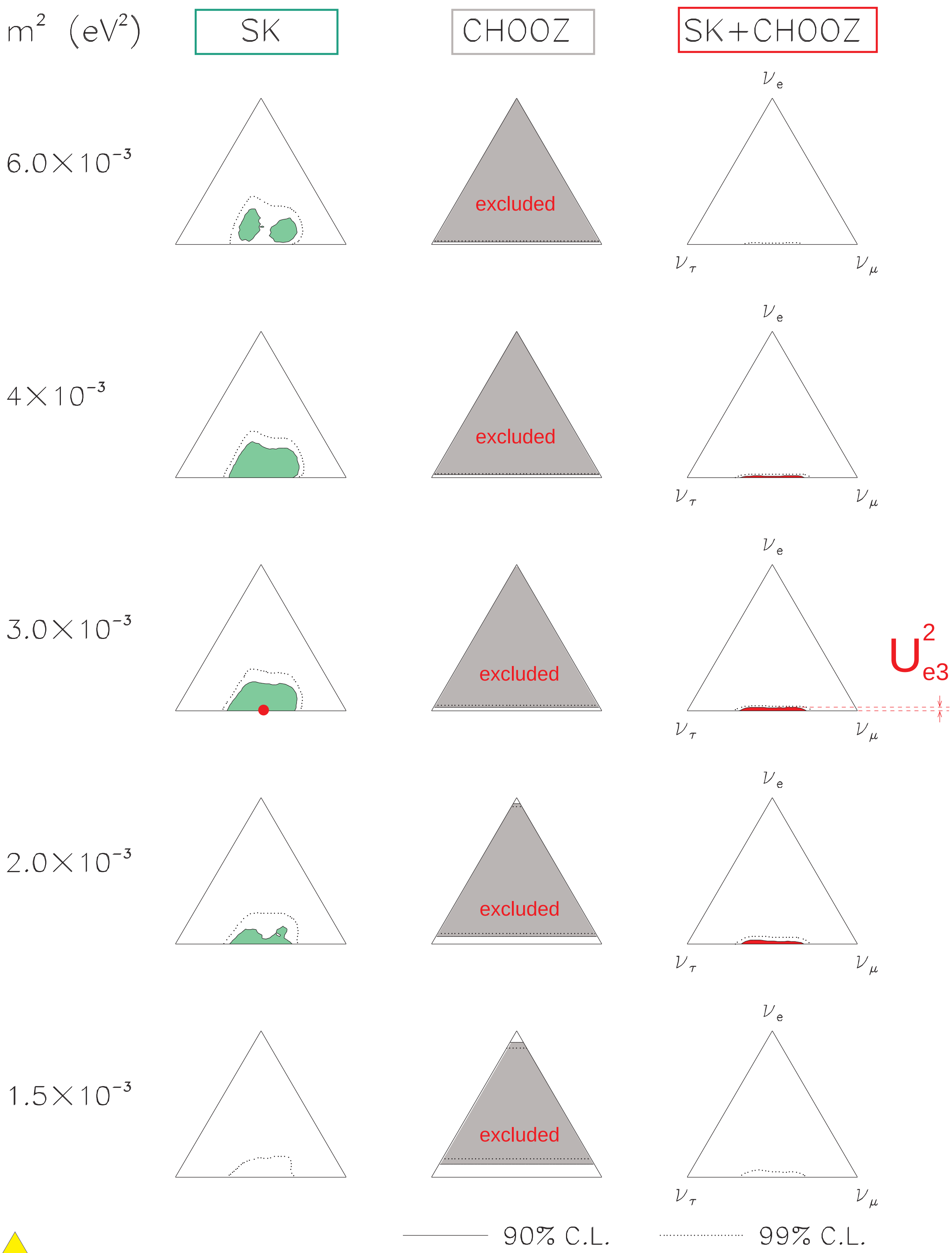
Allowed regions in a three-flavour approach

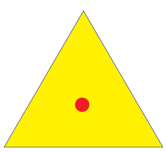
(Dec. 2000 SK data: 79.5 kTy)



Combining Superkamiokande and CHOOZ

(Dec. 2000 SK data: 79.5 kTy)

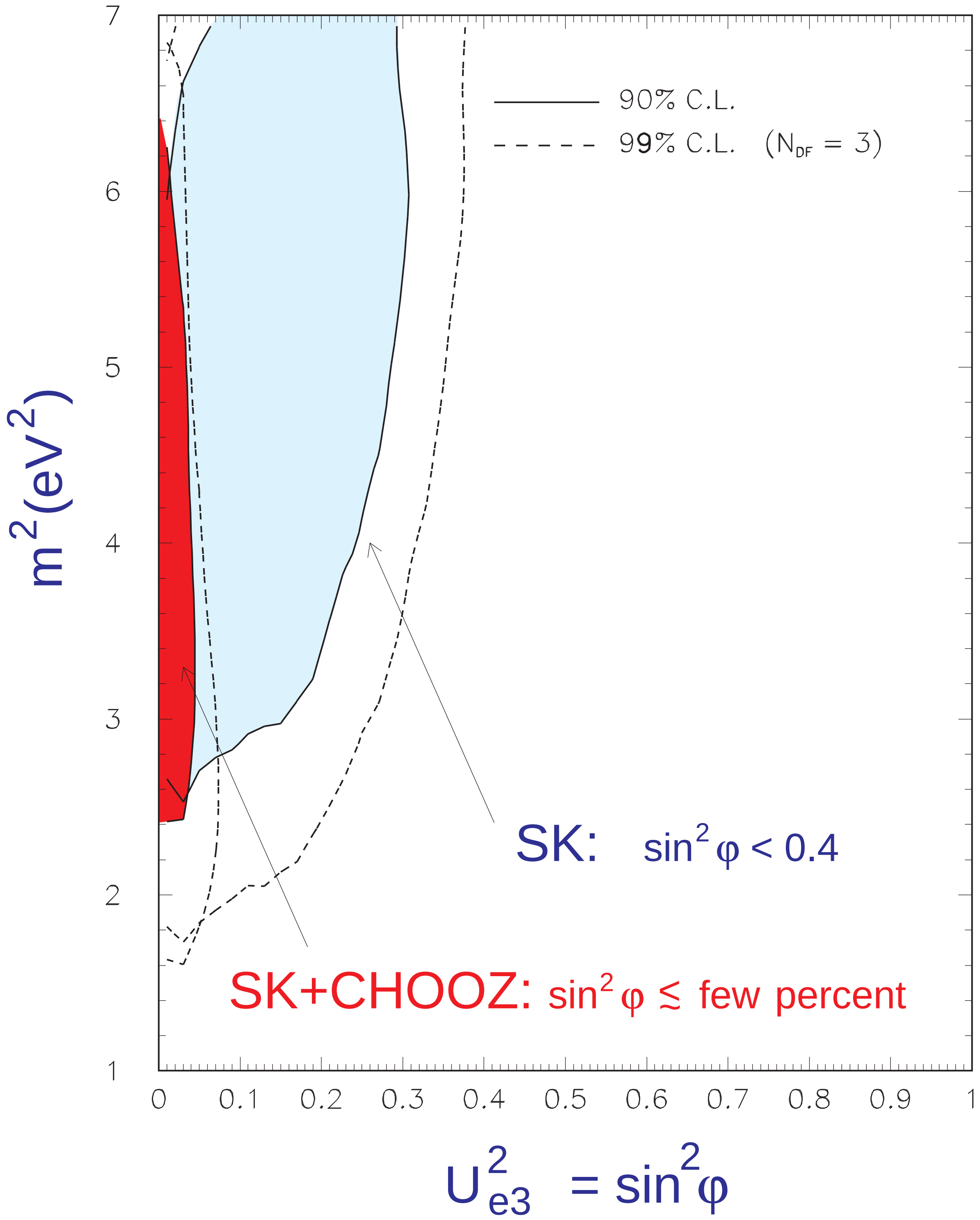




scenarios with large ν_e mixing excluded, e.g. threefold maximal mixing

Limits on U_{e3}^2 (ψ unconstrained)

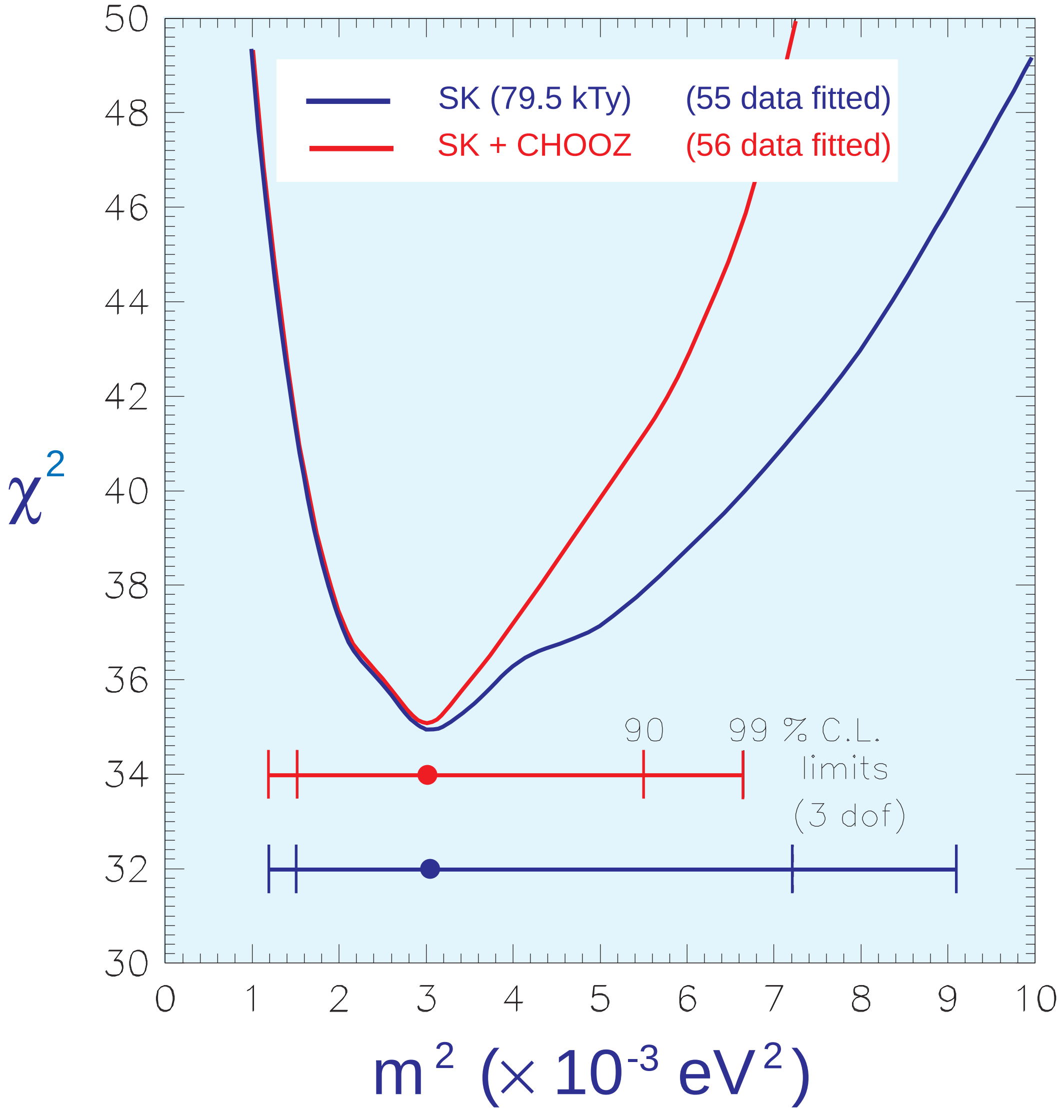
shown are the projections of the 3ν allowed volume onto the plane (m^2 , U_{e3}^2)



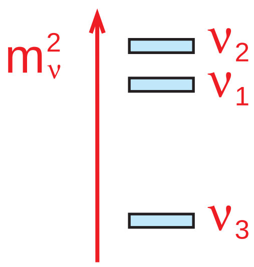
Note: threefold maximal mixing ($\sin^2\varphi = 1/3$) excluded by SK+CHOOZ, although allowed by SK alone

Bounds on m^2 for unconstrained 3ν mixing

(79.5 kTy SK data)



- best-fit @ $m^2 = 3.0 \times 10^{-3}$ eV^2

- 

The fit for $-m^2$ is very similar: SK does not distinguish the two cases. In the limit of pure $\nu_\mu \leftrightarrow \nu_\tau$, i.e. $\varphi = 0$,

$$P_{osc}(m^2) = P^{vac}_{osc}(m^2) = P^{vac}_{osc}(-m^2)$$

For small φ there are small differences due to matter effects, unobserved at present

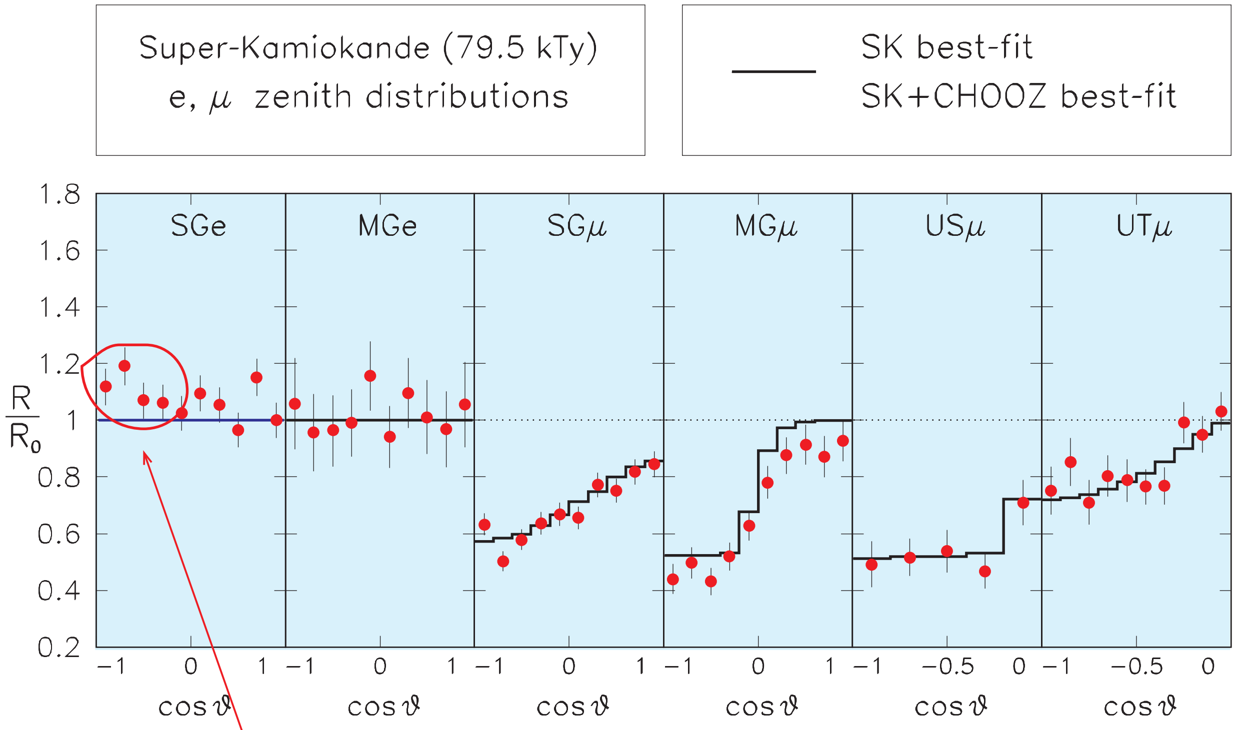
Best-fit distributions

SK alone

SK+CHOOZ

U^2_{e3} consistent with zero

(pure $\nu_\mu \leftrightarrow \nu_\tau$ favoured)



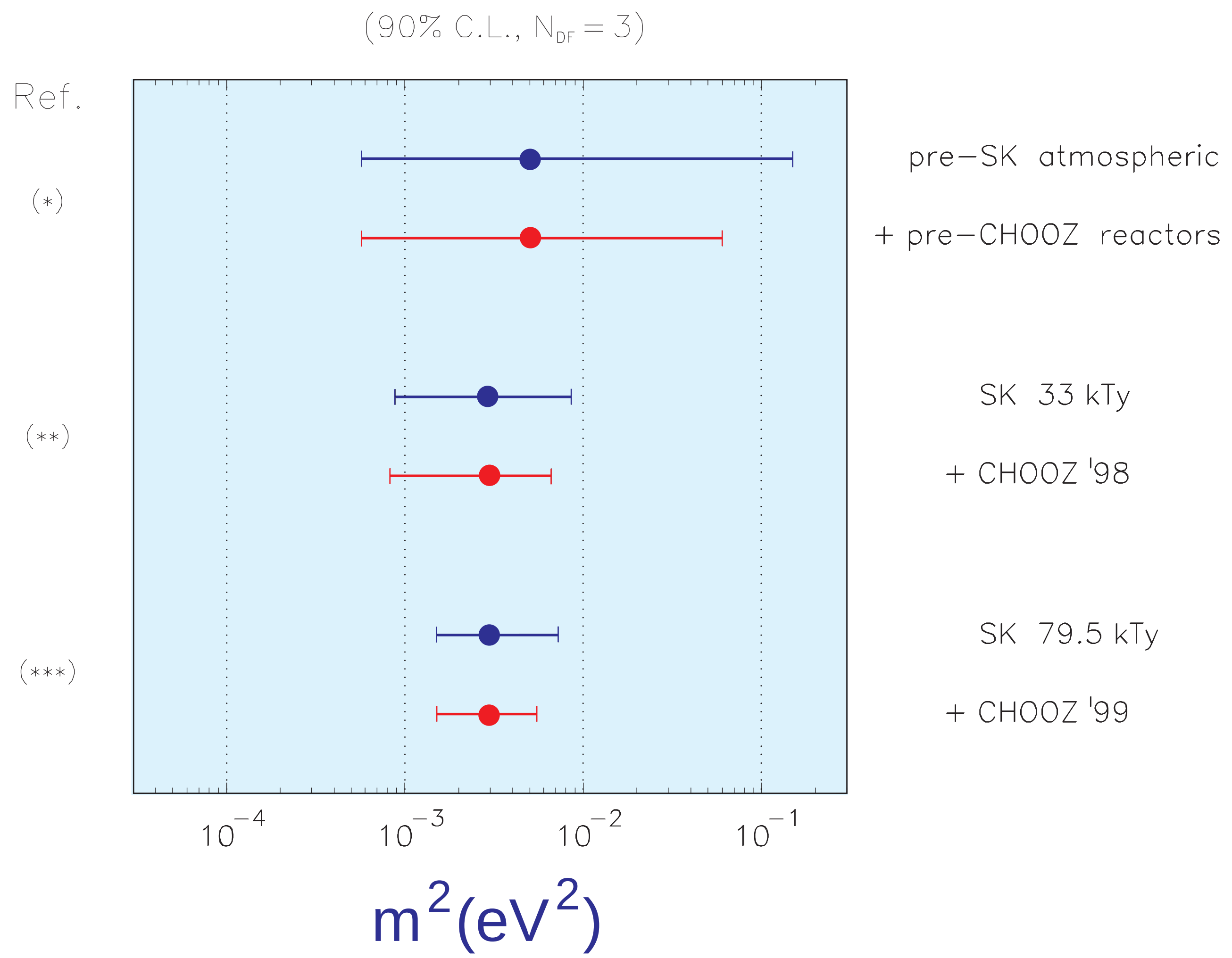
Possible deviations from flat due to $\nu_\mu \leftrightarrow \nu_e$ oscillations are typically smaller than these

Sources of possible deviations:

- $\varphi \neq 0$
- subleading δm^2 effects if $\delta m^2 \sim 10^{-4}$ eV^2

Both effects may alter the SGeV/MGeV e distributions: much higher SK statistics needed to see such effects !!

Progress in m^2 bounds for unconstrained 3ν mixing



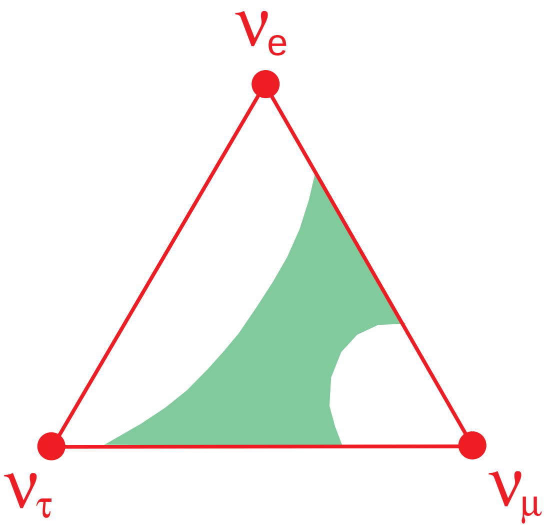
(*) GLF, E. Lisi, D. Montanino and G. Scioscia, PRD 55 (1997) 4385

(**) GLF, E. Lisi, A. Marrone and G. Scioscia, PRD 59 (1999) 033001

(***) GLF, E. Lisi and A. Marrone, this talk

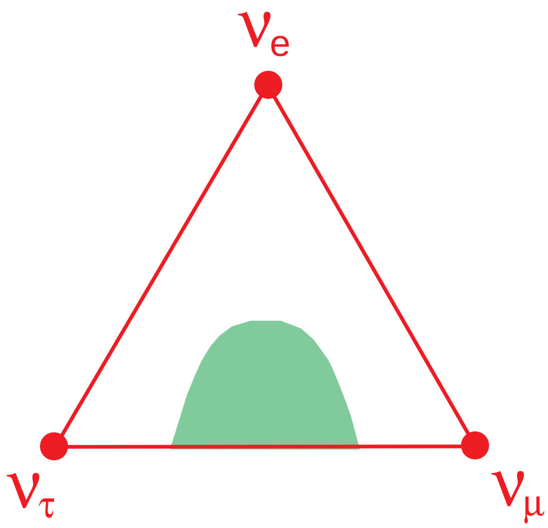
Tremendous impact of Superkamiokande in constraining m^2

Progress in constraining mixing (qualitative)



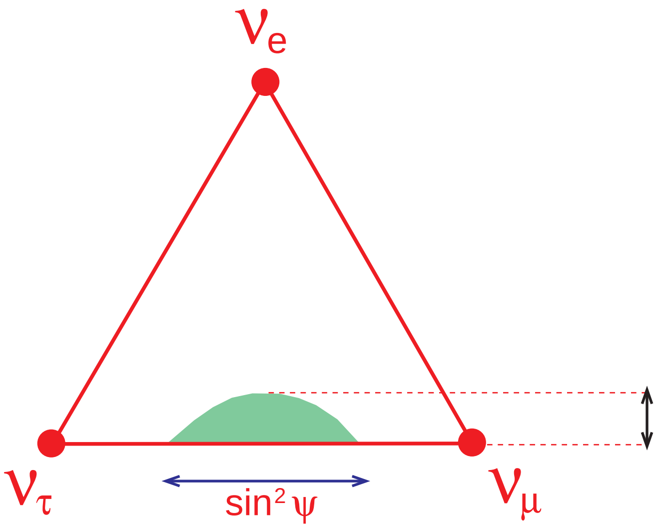
pre-SK and pre-CHOOZ data

both $\nu_\mu \leftrightarrow \nu_\tau$, $\nu_\mu \leftrightarrow \nu_e$ and their 3ν interpolations allowed



after SK

$\nu_\mu \leftrightarrow \nu_e$ excluded



after CHOOZ

$U_{e3}^2 = \sin^2\varphi$

Tremendous impact of CHOOZ in constraining φ

Status of (m^2, ψ, φ) constraints

1. $m^2 \simeq 3 \times 10^{-3}$ eV^2 within a factor of 2 ($1.5 \div 6.0 \times 10^{-3}$ eV^2)

2. $s^2_\psi \simeq 0.5 \pm 0.17$

3. $s^2_\varphi \lesssim$ few percent

Prospects

- SK will steadily narrow the range of (m^2, s^2_ψ) until systematics will dominate. Next major improvement will be provided by LBL experiments.

- Signals of $s^2_\varphi \neq 0$ are, and will be, more difficult to observe:

SK $\Rightarrow$ Typical signals of $s^2_\varphi \neq 0$ are smaller than present 1σ statistical uncertainties. To establish them at the 2σ level, more than 4 years are required....(much more if systematics are included ...)

LBL $\Rightarrow$ Signals of $s^2_\varphi \neq 0$ should be searched in the $\nu_\mu \leftrightarrow \nu_e$ channel, (with $P_{e\mu} \propto s^2_\varphi$). However, CHOOZ implies that $S/B \lesssim 1$ in LBL. So, the e-flavor background should be known precisely. This is difficult but important: $\varphi \neq 0$ is the only chance to observe MSW effect with terrestrial exps. (apart from the "exotic" $\nu_\mu \leftrightarrow \nu_s$ case)

ν factories $\Rightarrow$ Of course, they may provide the real option to observe $s^2_\varphi \neq 0$ in a "relatively far" future

reactors $\Rightarrow$ Next logical step to increase the sensitivity to s^2_φ is to place a near detector. There is a proposal of Krasnoyarsk.

A remark on LBL

- If m^2 is not too low, it is likely that K2K can observe ν_μ disappearance and confirm SK (hints from current K2K data)

- Therefore, LBL experiments should concentrate on what cannot be done in SK and K2K or in new atmospheric ν detectors, namely:

$\Longrightarrow$ ν_τ appearance

$\Longrightarrow$ ν_e appearance

appearance searches should be the main distinctive feature of LBL experiments

N.B. - Recent (weak) evidence for ν_τ appearance in SK, but on a statistical basis: only LBL (e.g. OPERA) should assess ν_τ appearance on a event by event basis

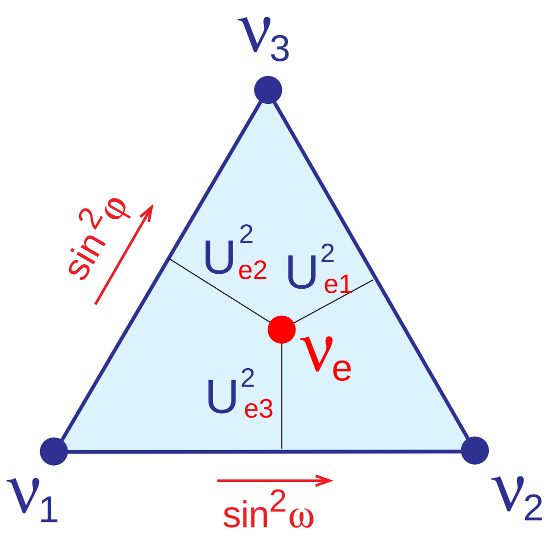
4. 3ν oscillations: solar neutrinos

$$\nu_e = U_{e1}\,\nu_1 + U_{e2}\,\nu_2 + U_{e3}\,\nu_3$$
$$= s_\phi\,\nu_3 + c_\phi\,(s_\omega \nu_1 + c_\omega \nu_2)$$

being

$$U_{e3}^2 = s_\phi^2$$
$$U_{e2}^2 = c_\phi^2\, c_\omega^2$$
$$U_{e1}^2 = c_\phi^2\, s_\omega^2$$

@ δm^2 fixed

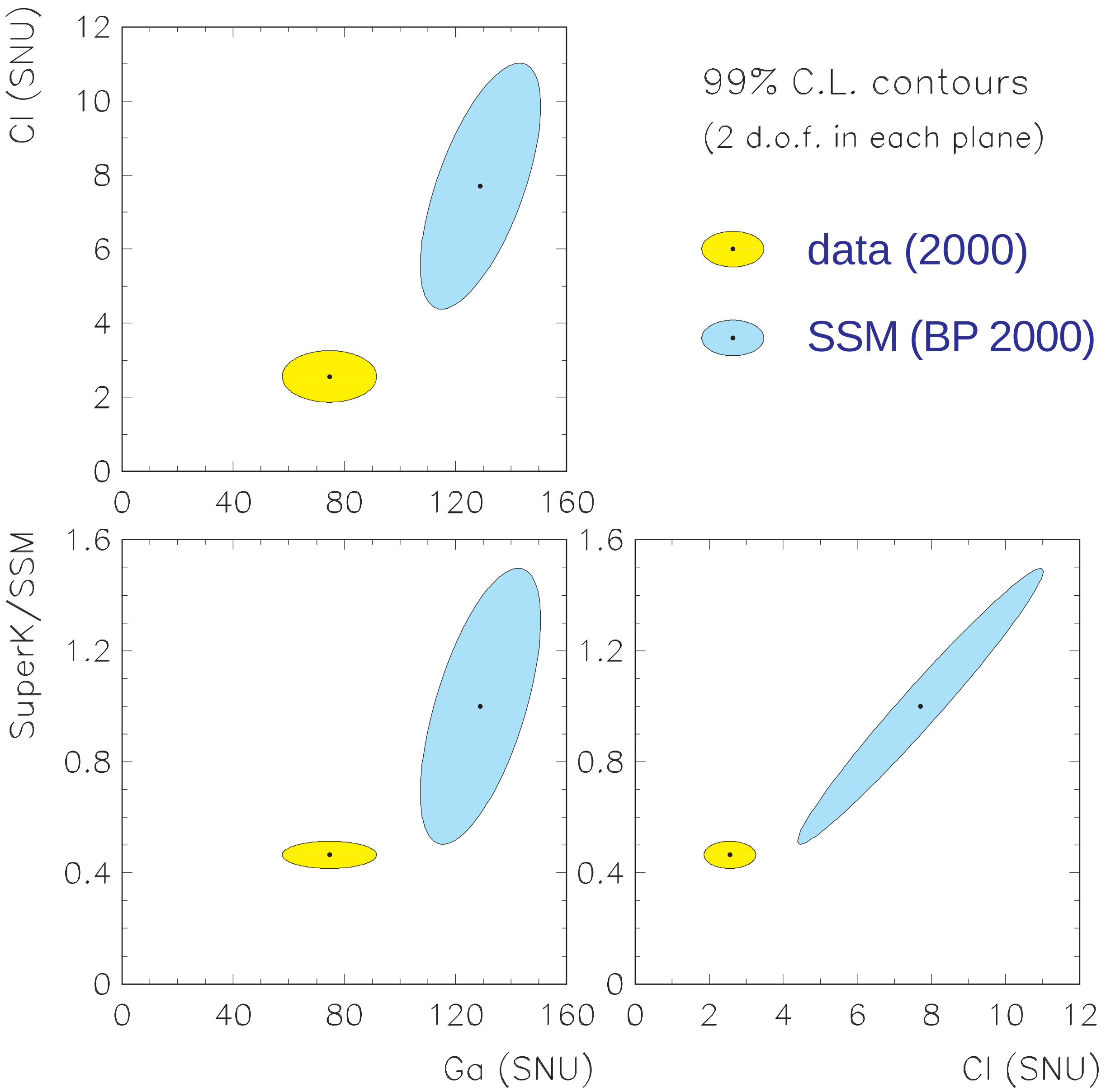


The analysis includes:

- 4 total rates: Cl = Homestake
  Ga = GALLEX+SAGE
  SK = SuperKamiokande (data Dec 2000)
- 18 SK energy bins (-1 free renormalization factor)
- day-night effect from SK including separately Sp(D) & Sp(N)

In the following we will assume: SSM = BP 2000
ϕ(hep) = SSM value

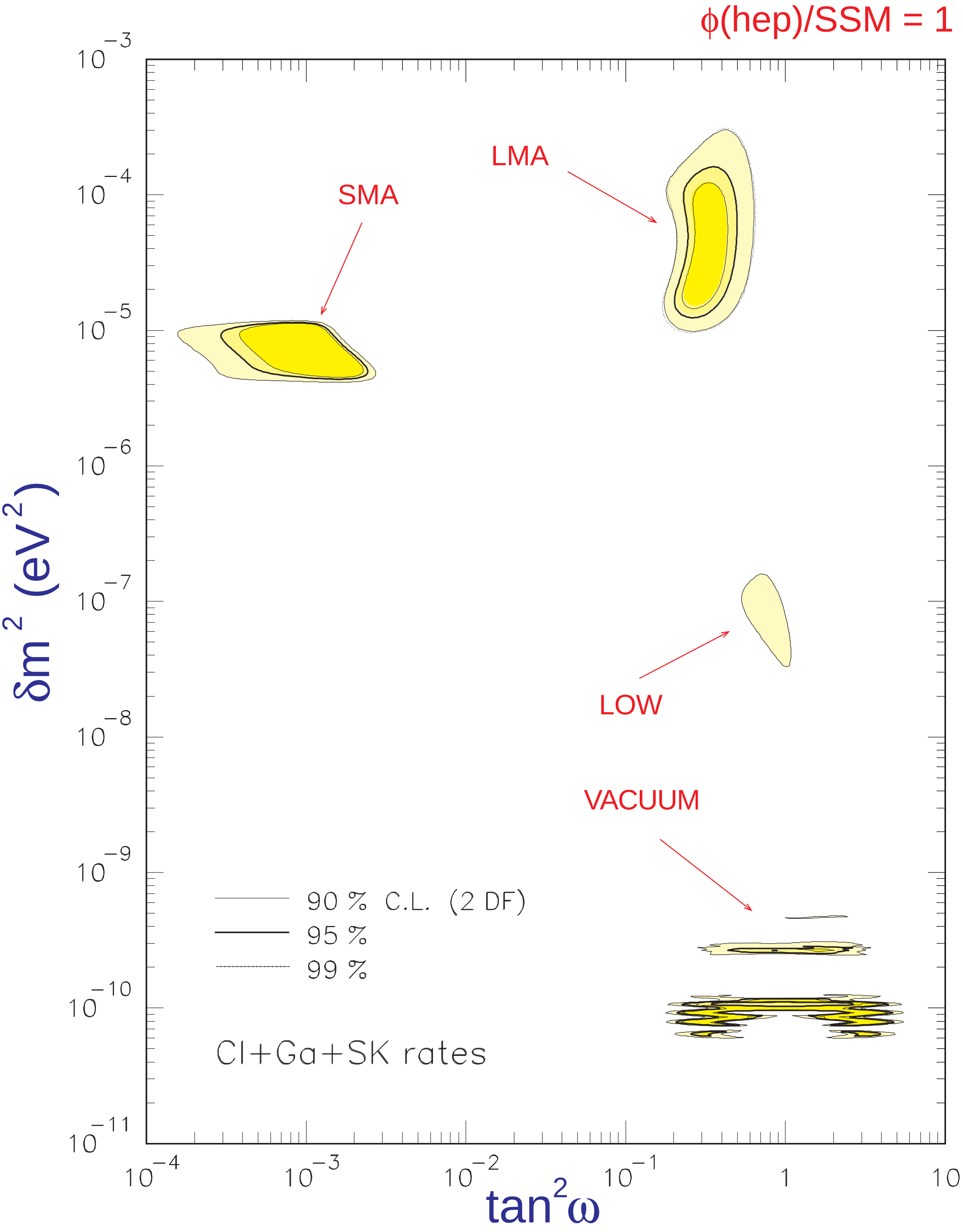
Solar neutrino problem, 2000



- clear evidence of the solar neutrino deficit!
- theor. error >> expt. error

We do not expect to improve our understanding of the solar ν deficit by just increasing the accuracy of the total flux measurements

2ν oscillations (φ = 0): total rates



SMA: $\chi^2 \sim 0.25$ best fit (2 dof)

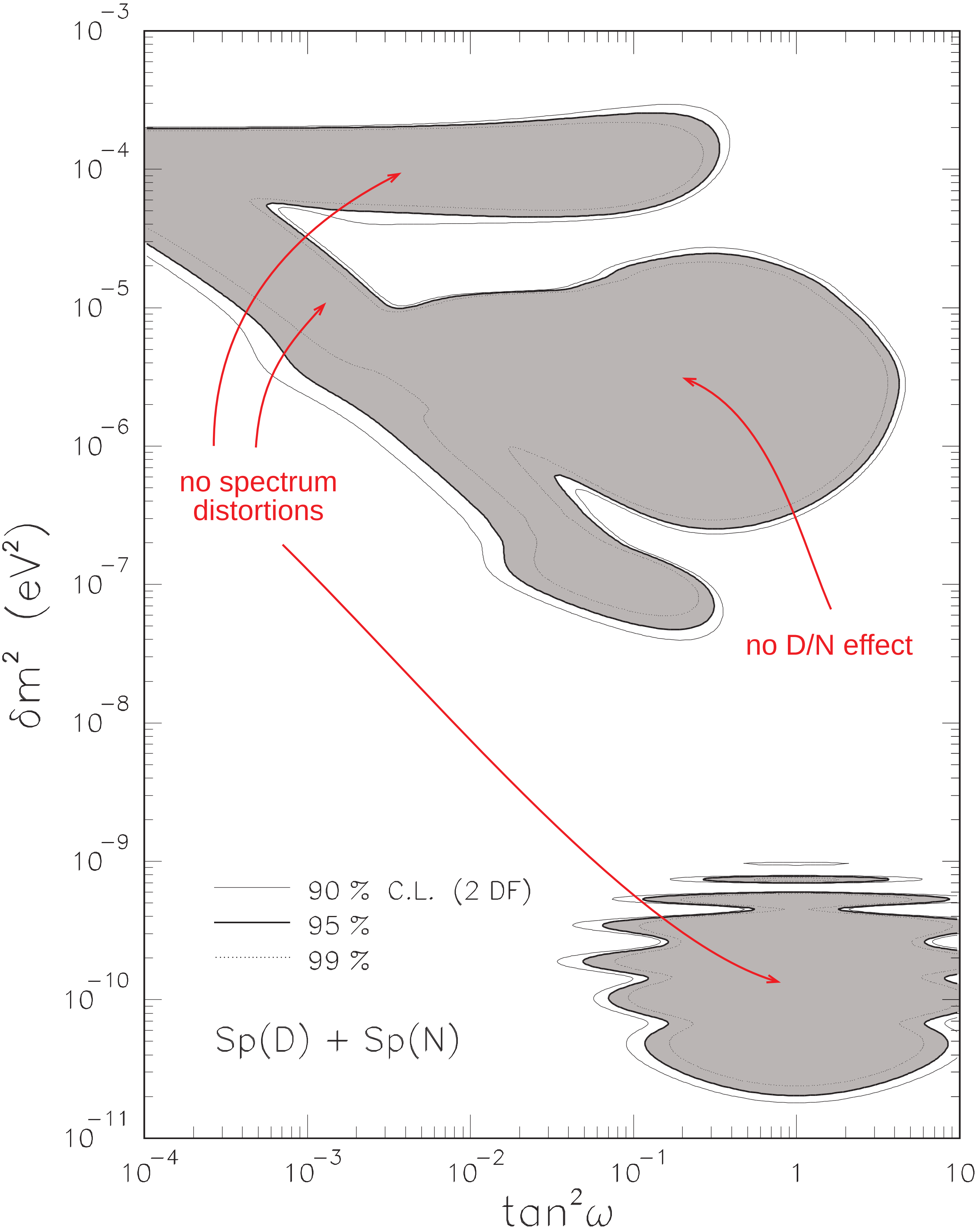
VACUUM: $\chi^2 \sim 0.26$

LMA: $\chi^2 \sim 3.55$

LOW: $\chi^2 \sim 8.05$

2ν oscillations (φ = 0): SK spectrum

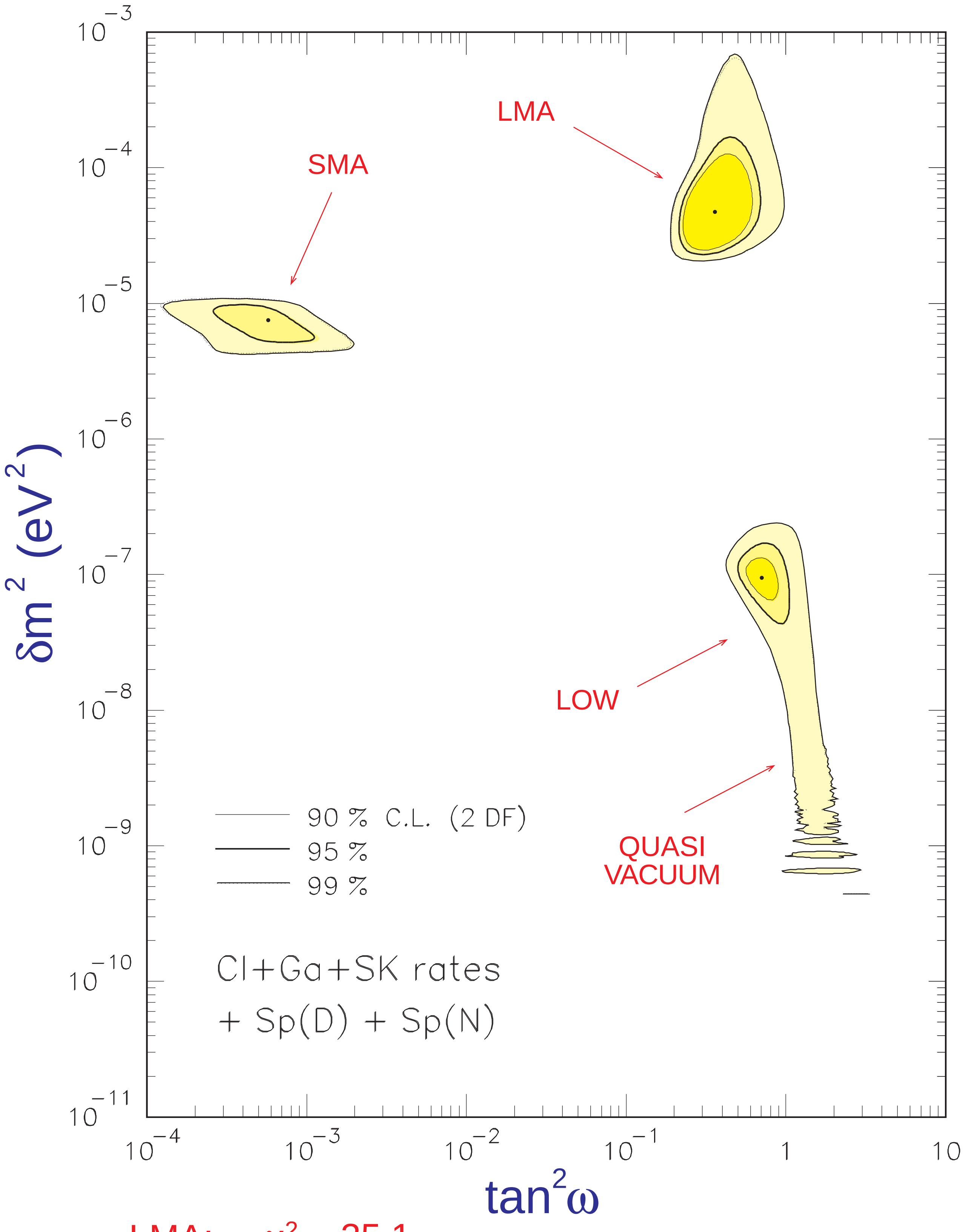
regions excluded by Sp(D) + Sp(N) (hep/SSM=1)



Comment: ● Poor overlap between regions allowed by spectrum and by rates (see the previous figure)

2ν oscillations (φ = 0)

total rates and SK spectrum (N_{dof} = 36)

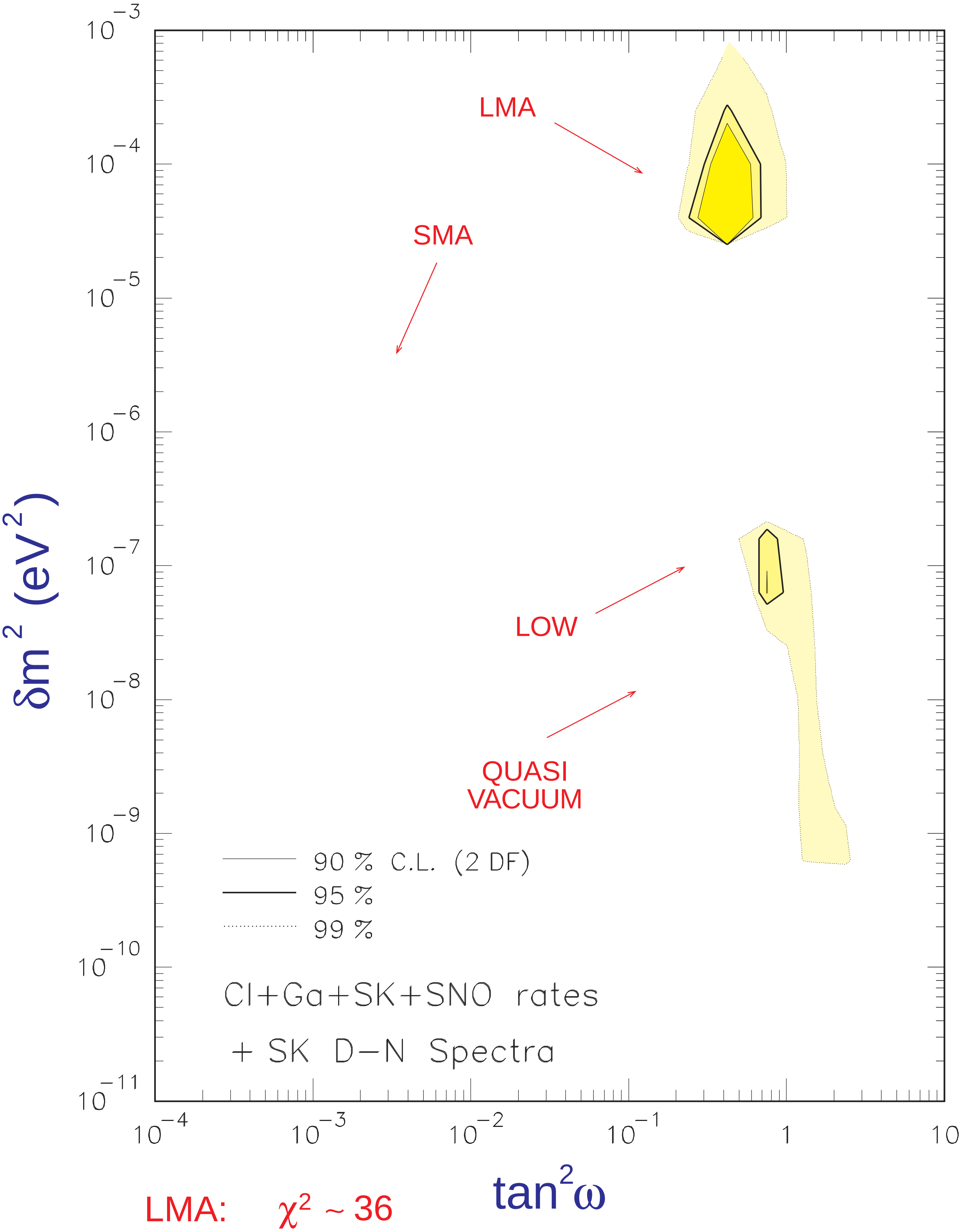


LMA: $\chi^2 \sim 35.1$
LOW: $\chi^2 \sim 39.0$
SMA: $\chi^2 \sim 40.8$

More spectral data needed to prefer or to exclude one of the solutions

2ν oscillations ($\varphi = 0$)

Inclusion of the SNO data !!



LMA: $\chi^2 \sim 36$
LOW: $\chi^2 \sim 40$
SMA: $\chi^2 \sim 49$

Towards a 3ν analysis

2ν → 3ν

$(\delta m^2, \omega) \xrightarrow[m^2 \to \infty]{} (\delta m^2, \omega, \varphi)$

$P_{2\nu}(\nu_e \to \nu_e) \xrightarrow[m^2 \to \infty]{} P_{3\nu} = c_\varphi^4 P_{2\nu}\Big|_{N_e = c_\varphi^2 N_e} + s_\varphi^4$

φ small (CHOOZ) implies that $P_{3\nu} \sim P_{2\nu}$, so why we study the case of unconstrained φ ?

Two reasons:

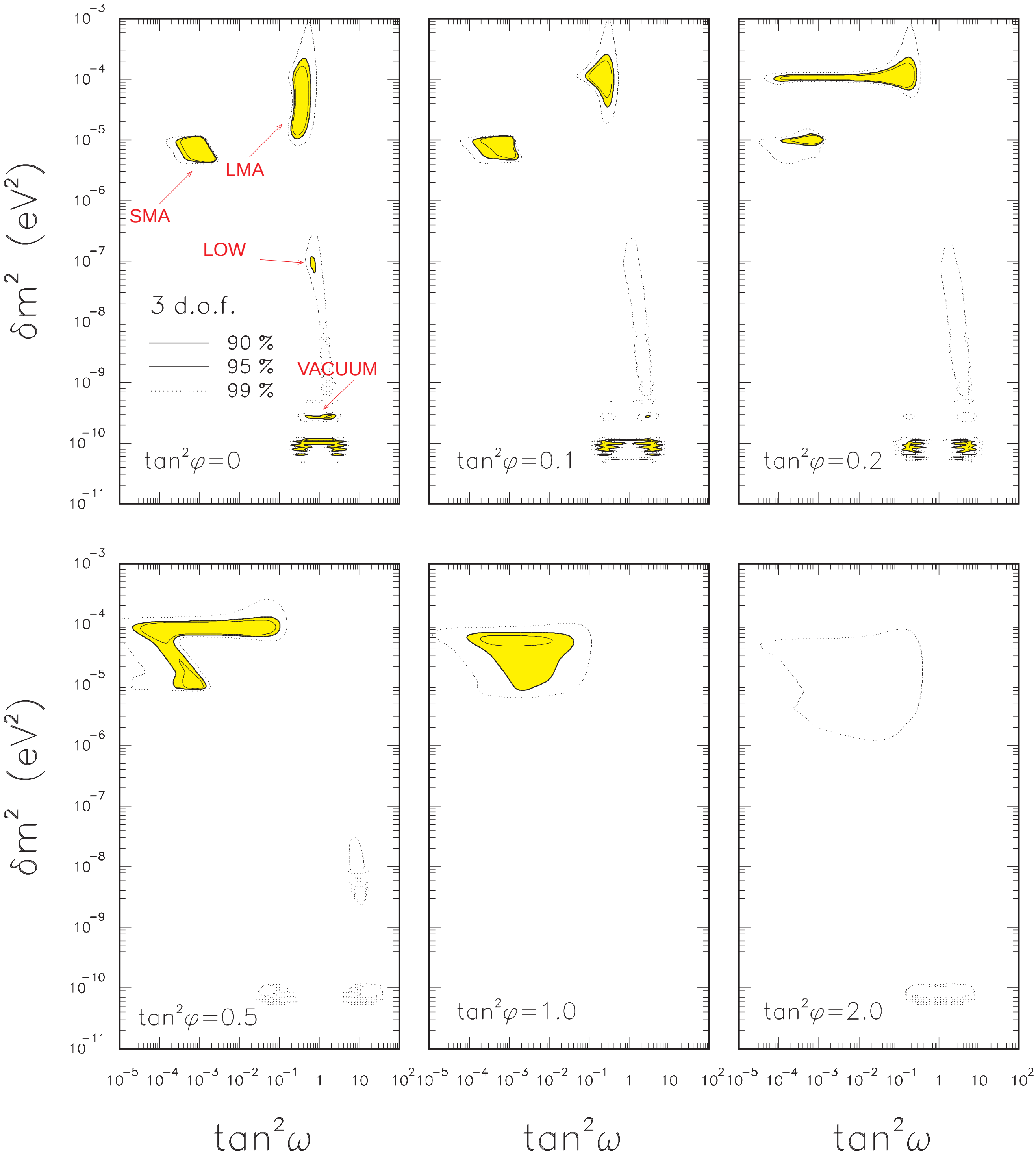
- Investigate if solar ν data alone (without CHOOZ) prefer small φ, in the same way as atmospheric data alone

- Study the behaviour of the usual 2ν solutions, in particular SMA, LMA and LOW, under small φ perturbations

In the following ϕ(hep) = SSM value is assumed.

3ν solutions

total rates from Cl+ Ga+K+SK (hep/SSM = 1)

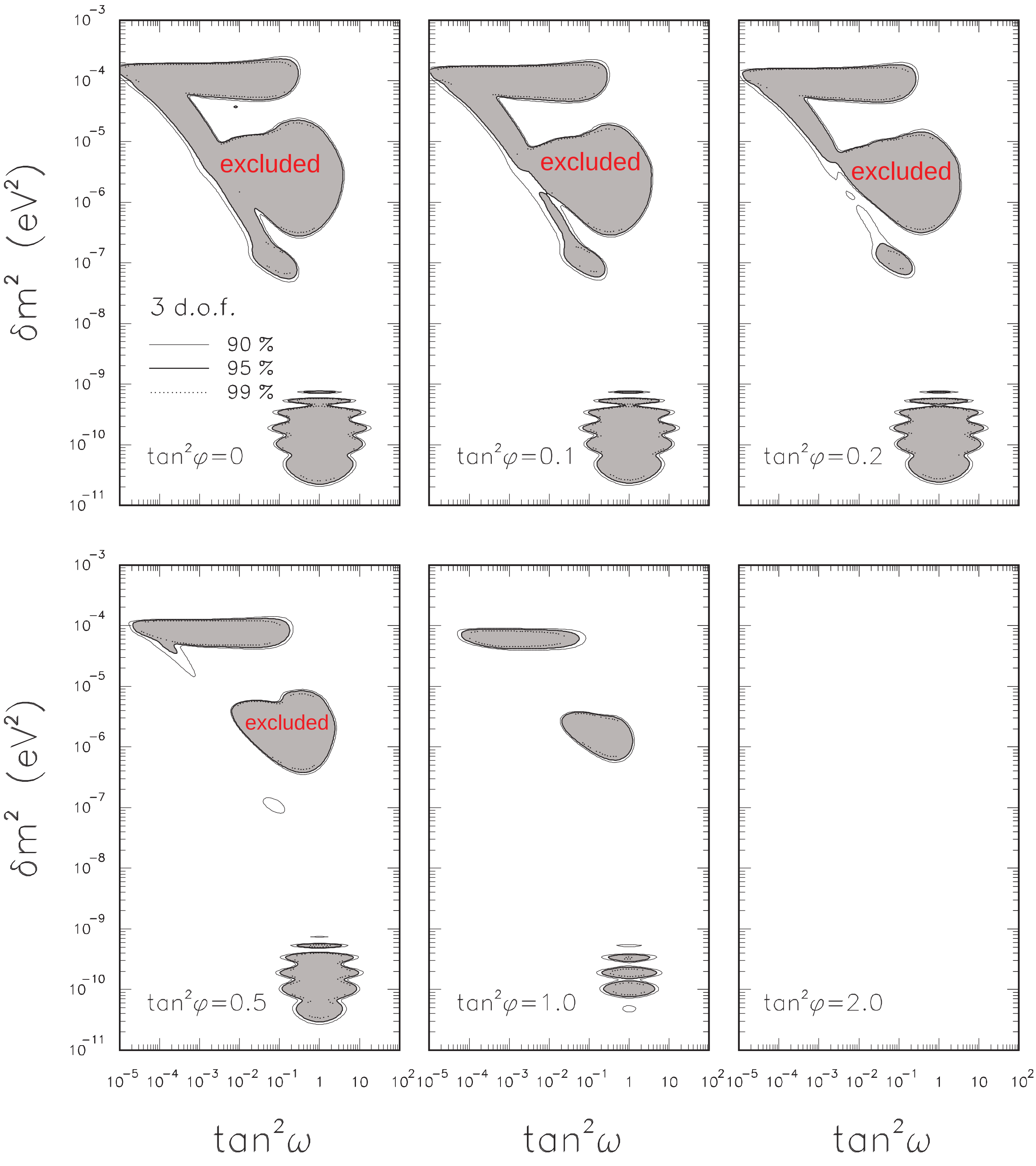


Comments:

- Best fit at $s^2_\varphi = 0$
- Loose upper bound on s^2_φ: $s^2_\varphi < 0.7$

3ν solution

constraints from Sp(Day) + Sp(Night) (hep/SSM=1)

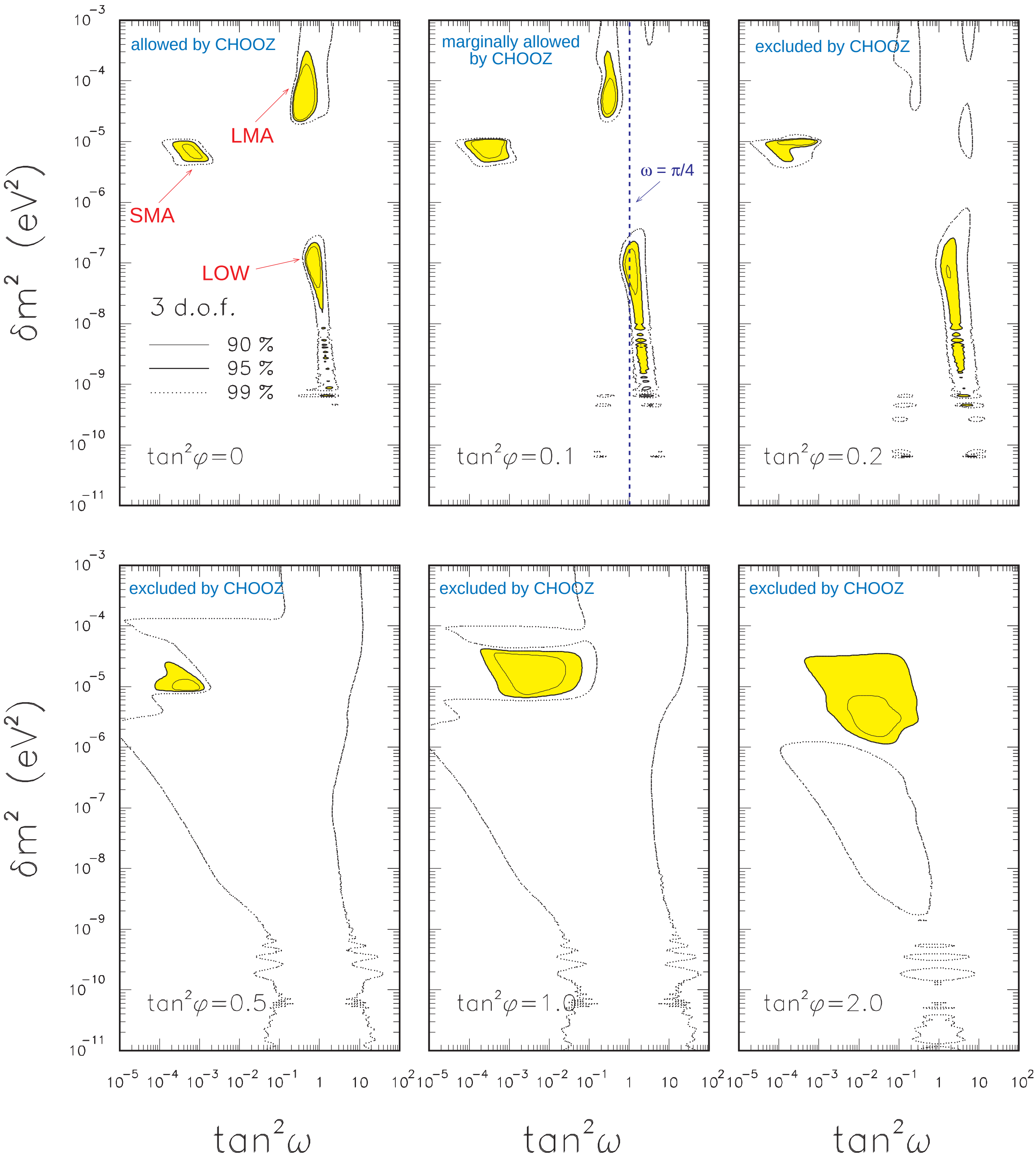


Comments: ● Sp(Day)+Sp(Night) consistent with no oscillation (s^2_φ = 1)

⇒ no bound on s^2_φ, although s^2_φ= 0 slightly preferred

3ν solutions (hep/SSM = 1)

total rates with constraints from Sp(D)+Sp(N)

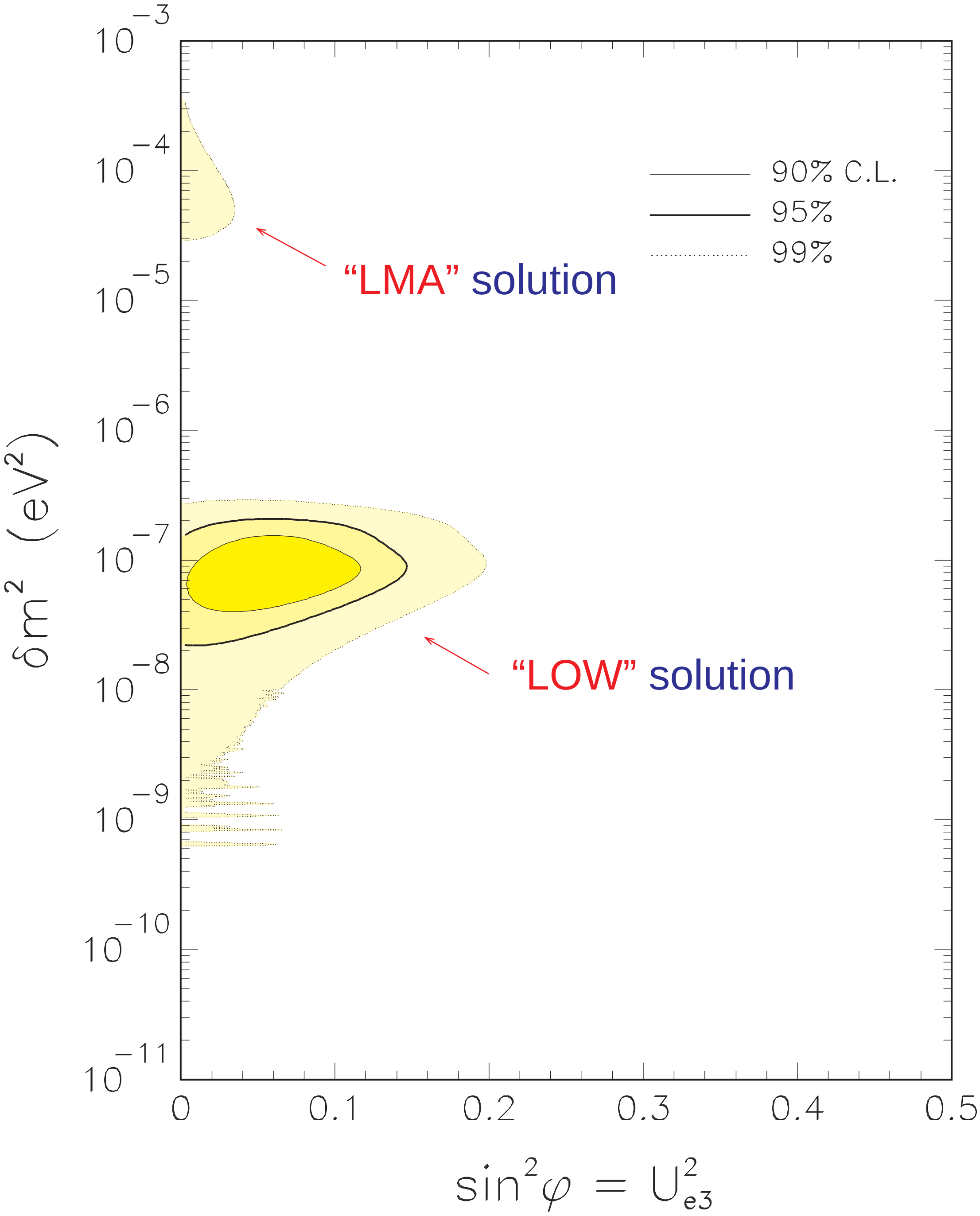


Comments:

- Still loose bounds ($s^2_\varphi < 0.7$) with $s^2_\varphi \sim 0.1$ preferred
- For small φ, maximal mixing solutions are allowed ! The LOW solution migrates toward $\omega = \pi/4$, and one can even have solutions for $\omega = \pi/4 + \varepsilon$ (completely missed if one uses $\sin^2 2\omega$ as variable)

3ν solutions @ maximal mixing

Comparison between "LOW" and "LMA" solutions
assuming maximal $\nu_{1,2}$ mixing: $U^2_{e1} = U^2_{e2}$ ($\omega = \pi/4$)



At maximal (ν_1, ν_2) mixing:

- the LOW solution is enhanced for $U^2_{e3} \sim 0.05$
- $\chi^2_{min} = 39.9$ (for $\delta m^2 = 7.8 \times 10^{-8}$ eV^2 and $\sin^2\varphi \sim 6 \times 10^{-2}$)

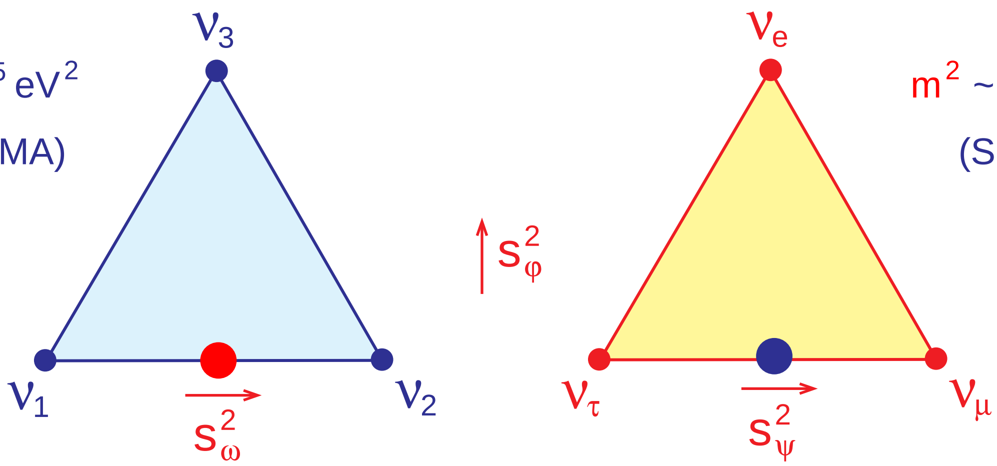
The previous result is interesting for model building: bimaximal mixing can be reached not only with the LMA solution, but also with the LOW solution, but as a bimaximal mixing at small φ:

1

Solar ν | Atmospheric ν

$\delta m^2 \sim 10^{-5}\ \mathrm{eV}^2$ (LMA)

$m^2 \sim 3 \times 10^{-3}\ \mathrm{eV}^2$ (SK + CHOOZ)



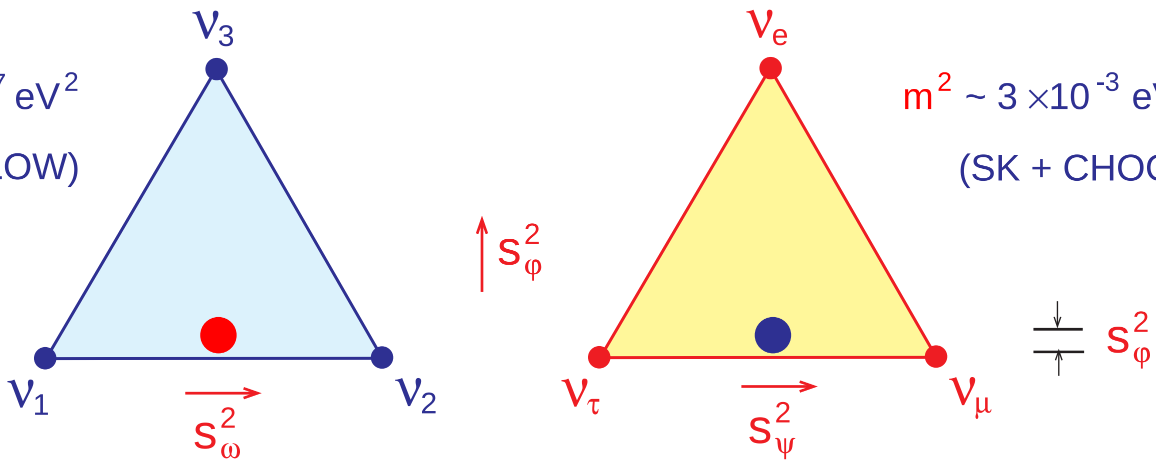
bimaximal mixing $\longrightarrow$ $\begin{cases} s^2_\omega = s^2_\psi = \frac{1}{2} \\ s^2_\varphi = 0 \end{cases}$

2

Solar ν | Atmospheric ν

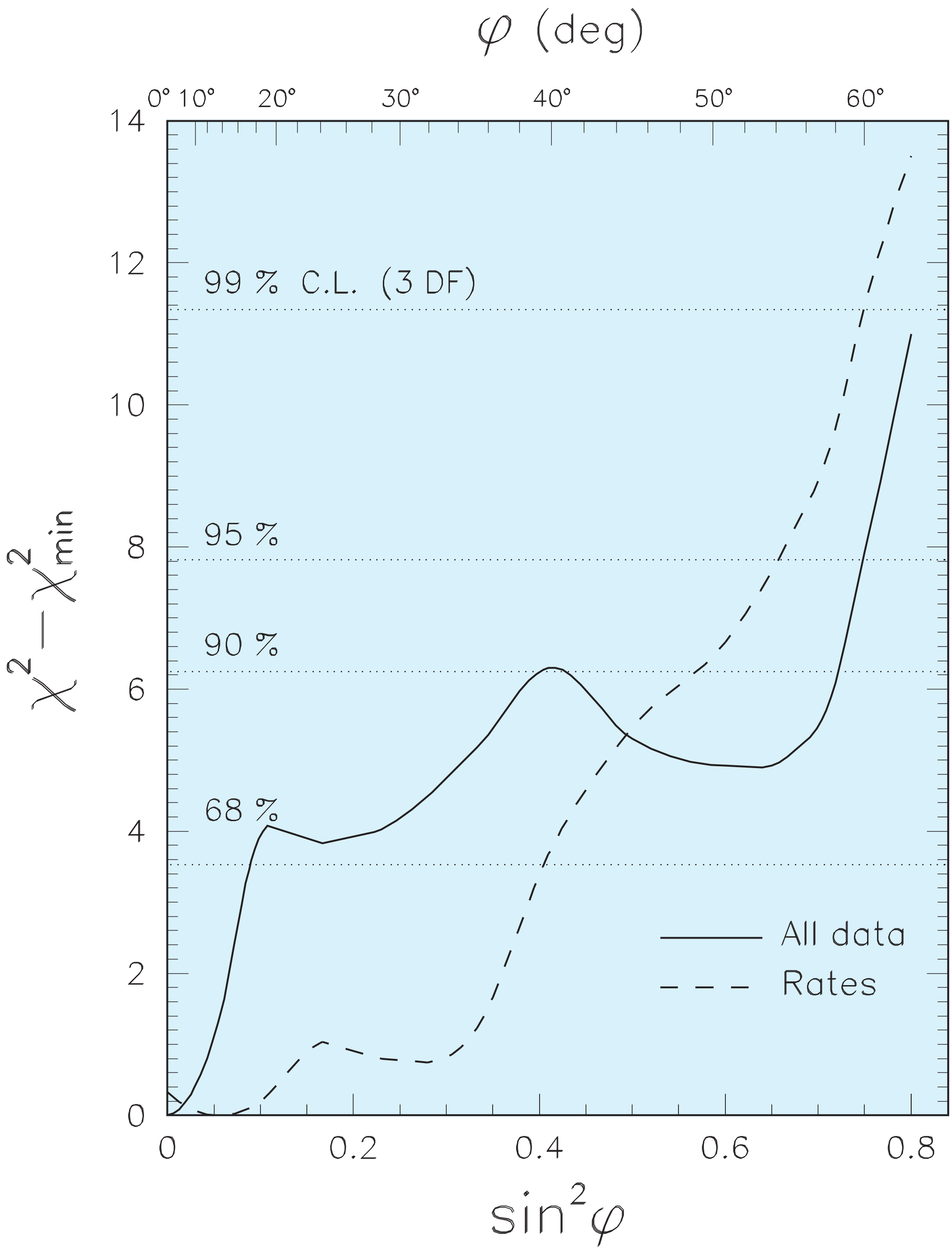
$\delta m^2 \sim 10^{-7}\ \mathrm{eV}^2$ (MSW LOW)

$m^2 \sim 3 \times 10^{-3}\ \mathrm{eV}^2$ (SK + CHOOZ)



bimaximal mixing at small φ $\longrightarrow$ $\begin{cases} s^2_\omega = s^2_\psi = \frac{1}{2} \\ s^2_\varphi \sim \text{few percent} \end{cases}$

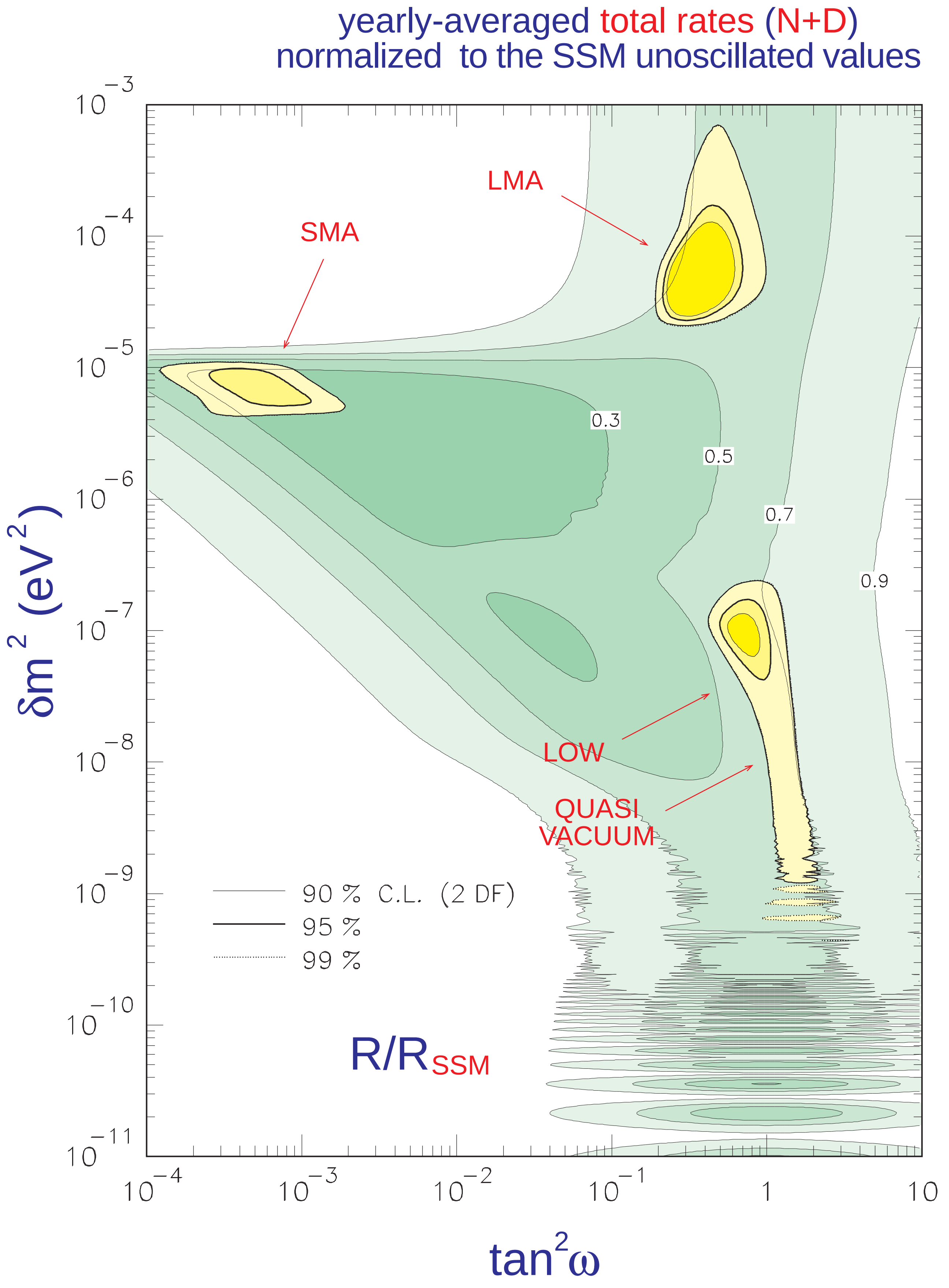
$\Delta\chi^2$ as a function of φ (δm^2 and ω unconstrained)



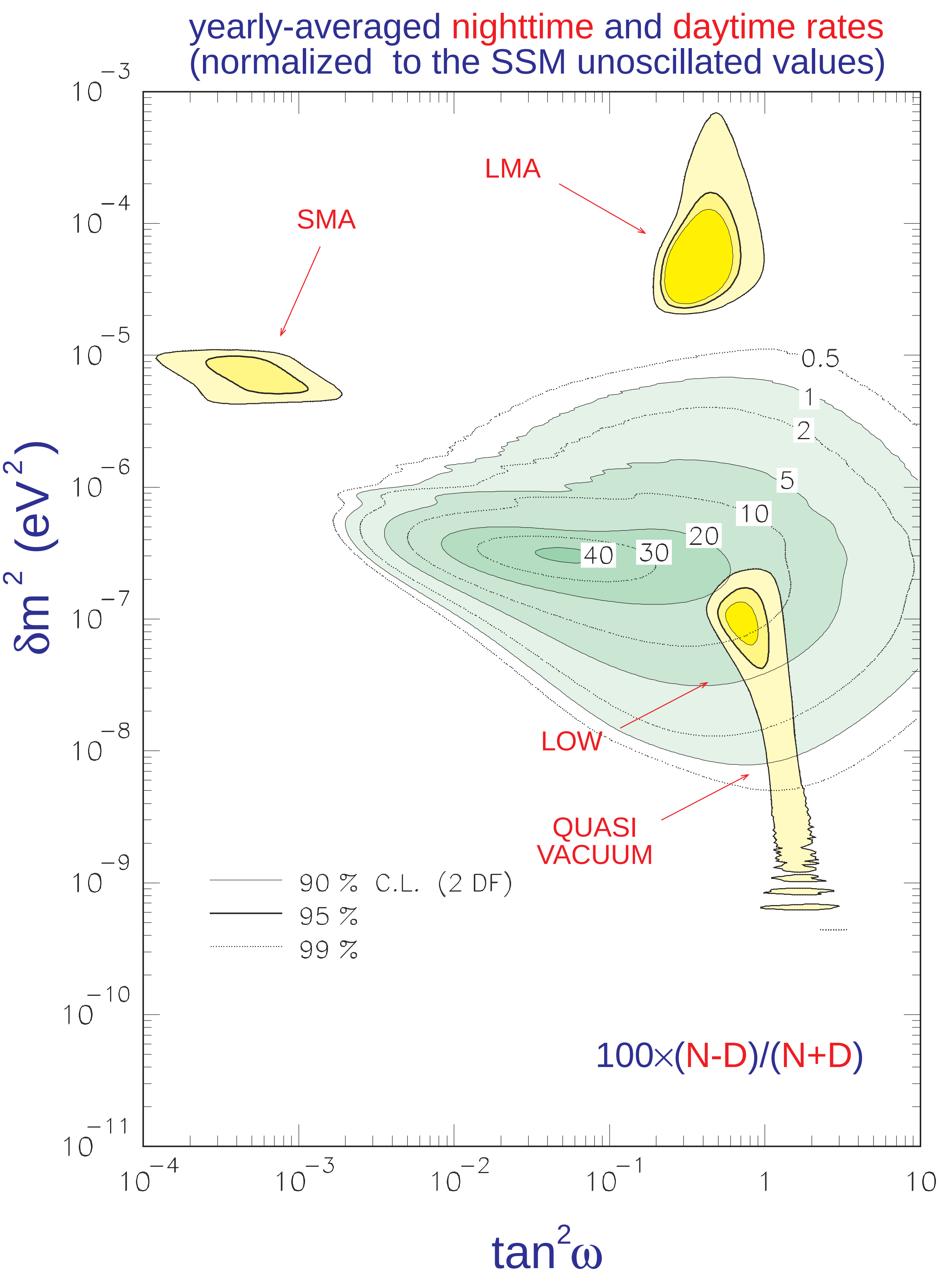
Comments:

- transition from the LMA to the LOW solution for $\sin^2\varphi$ larger than 0.1
- existence of an upper bound on $\tan^2\varphi$

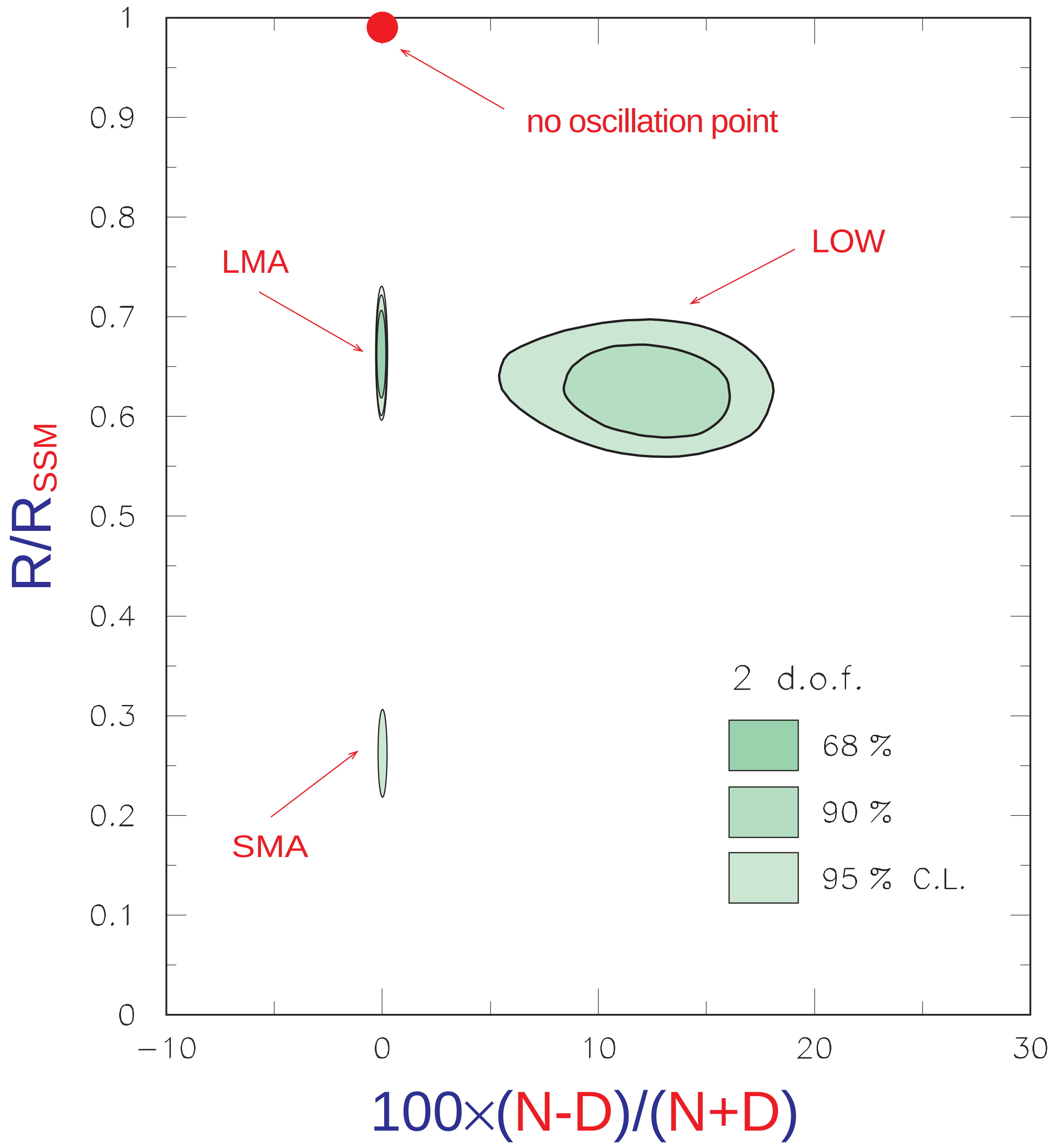
Borexino total rates compared with the SMA, LMA and LOW solutions



Borexino N-D asymmetry compared with the SMA, LMA and LOW solutions

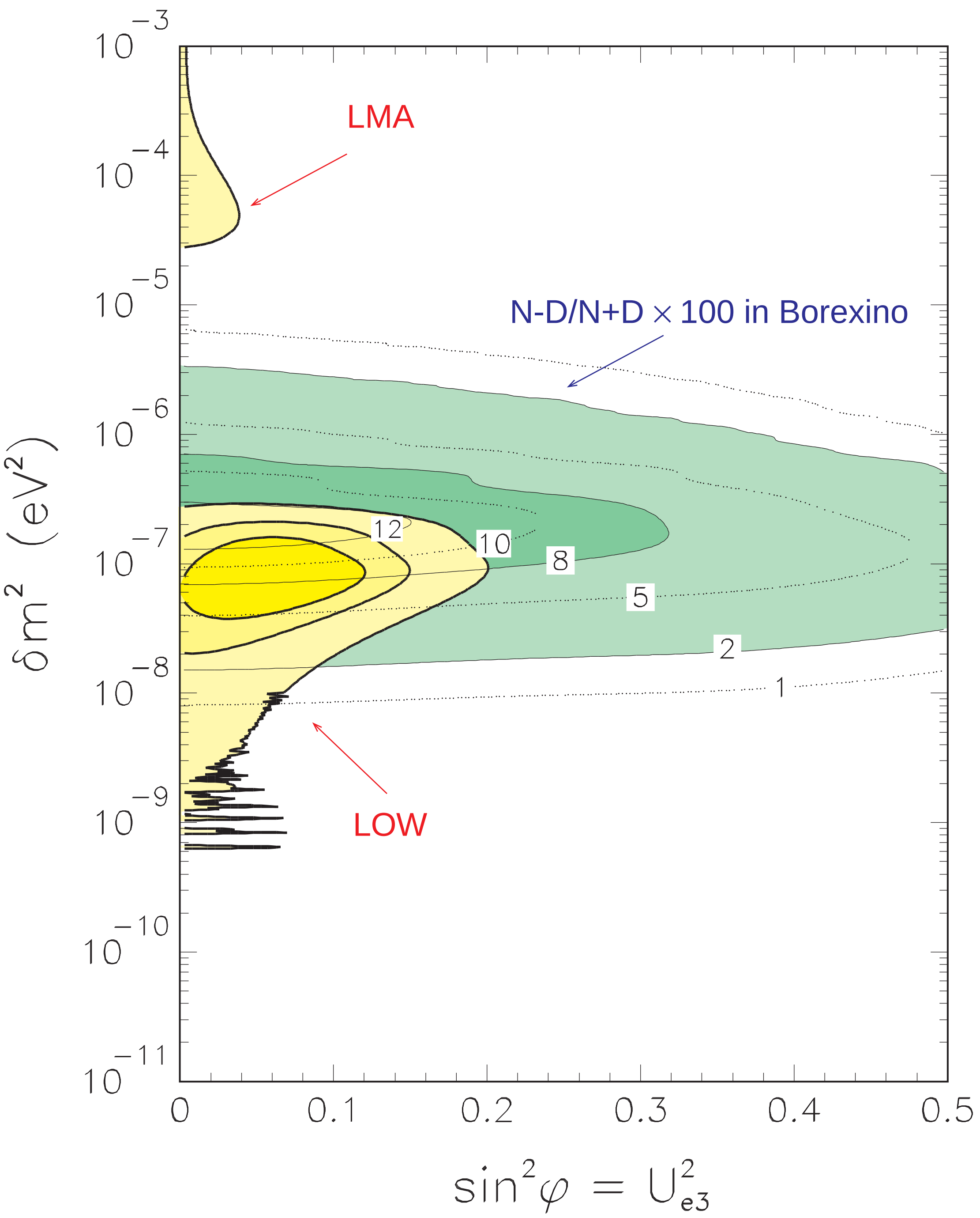


Borexino discovery potential compared with the SMA, LMA and LOW solutions



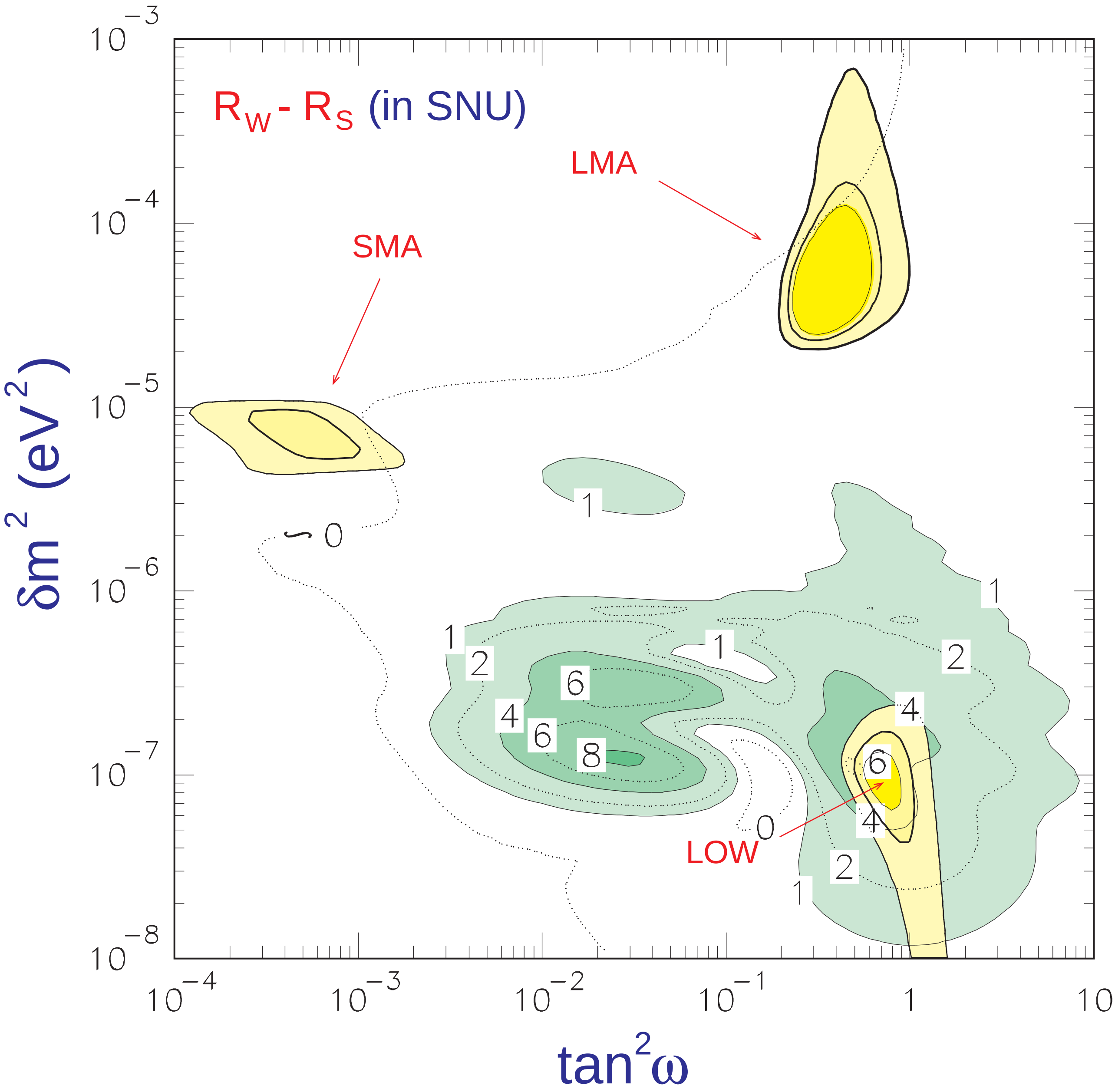
Borexino N-D asymmetry compared with the LMA and LOW solutions at maximal mixing (ω = π/4)

yearly-averaged nighttime and daytime rates
(normalized to the SSM unoscillated values)



Gallium Neutrino Observatory

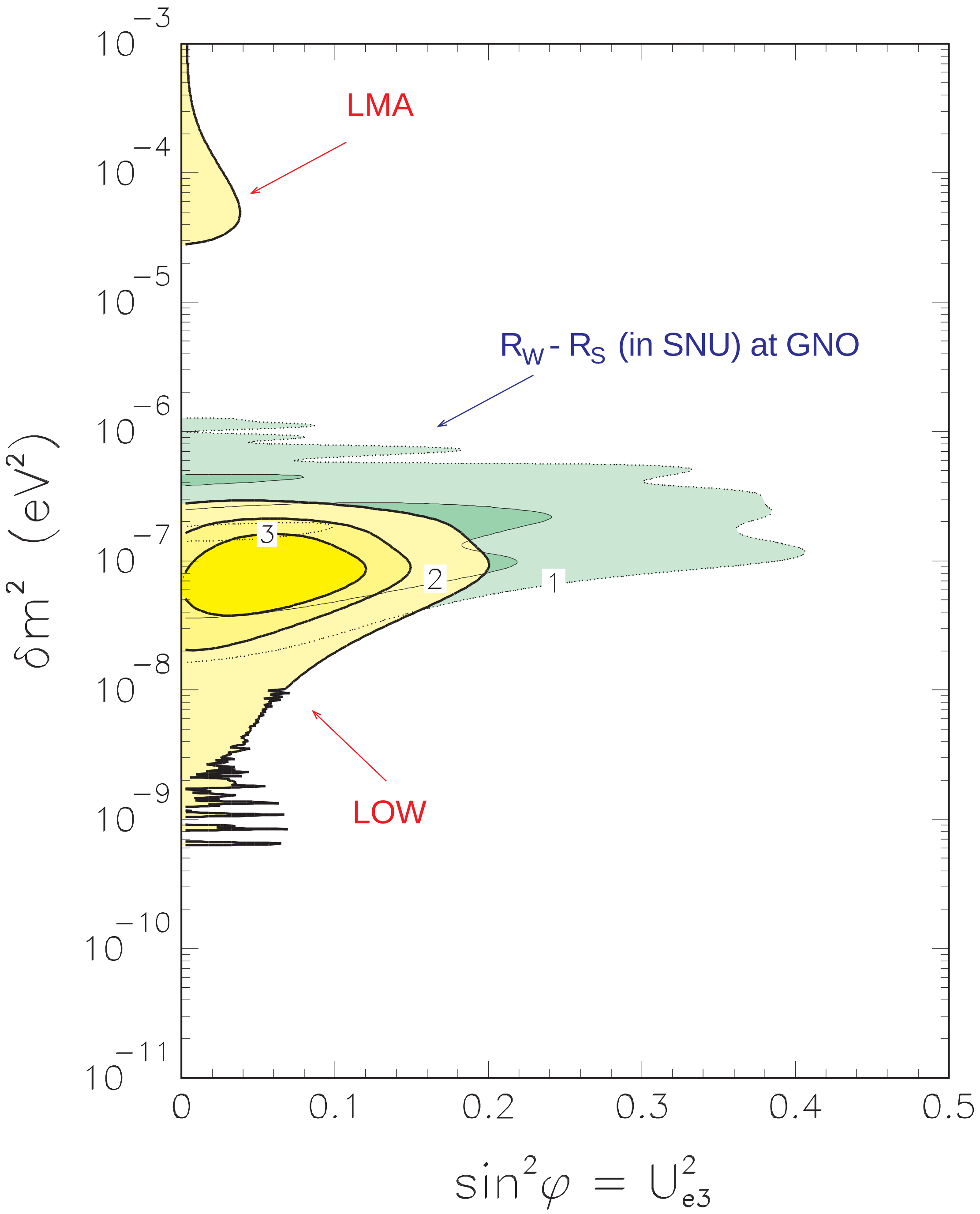
potential discovery of GNO compared with the SMA, LMA and LOW solutions



- expected gallium absorption rates (in SNU) averaged over "winter" and "summer" (eccentricity effects removed)
- MSW-induced seasonal variations of the order ~4-6 SNU are expected in a Ga experiment within the LOW solution (mainly from pp neutrinos)
- such variations might well be observed at GNO, the expected statistical error after one solar cycle being ~2 SNU (or less)

Gallium Neutrino Observatory

Comparison with the LMA and LOW solutions at maximal mixing ($\omega = \pi/4$)



- expected gallium absorption rates (in SNU) averaged over "winter" and "summer" (eccentricity effects removed)
- MSW-induced seasonal variations of the order ~4-6 SNU are expected in a Ga experiment within the LOW solution (mainly from pp neutrinos)
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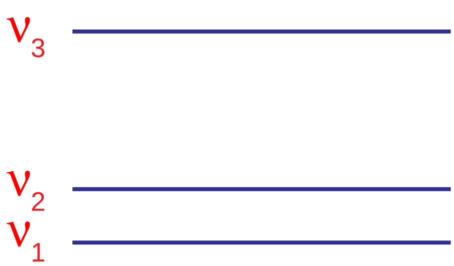
Beyond the “one dominant mass scale approximation”

$\Downarrow$

CHOOZ limits to the solar neutrino problem

1. CHOOZ excludes large ν_e mixing between any two states separated by $\Delta m^2 \geq 10^{-3}$ eV ^{2}

2. so, concerning m^2



since $m^2 \geq 10^{-3}$ eV ^{2}
(from atmospheric ν data)

$\Downarrow$

ν_3 must have small mixing with ν_e $\Rightarrow$ small $\langle \nu_3 | \nu_e \rangle = U_{e3} = s_\varphi$

3. and concerning δm^2

either ν_1 or ν_2 or both must be mixed with ν_e
(from solar ν data)

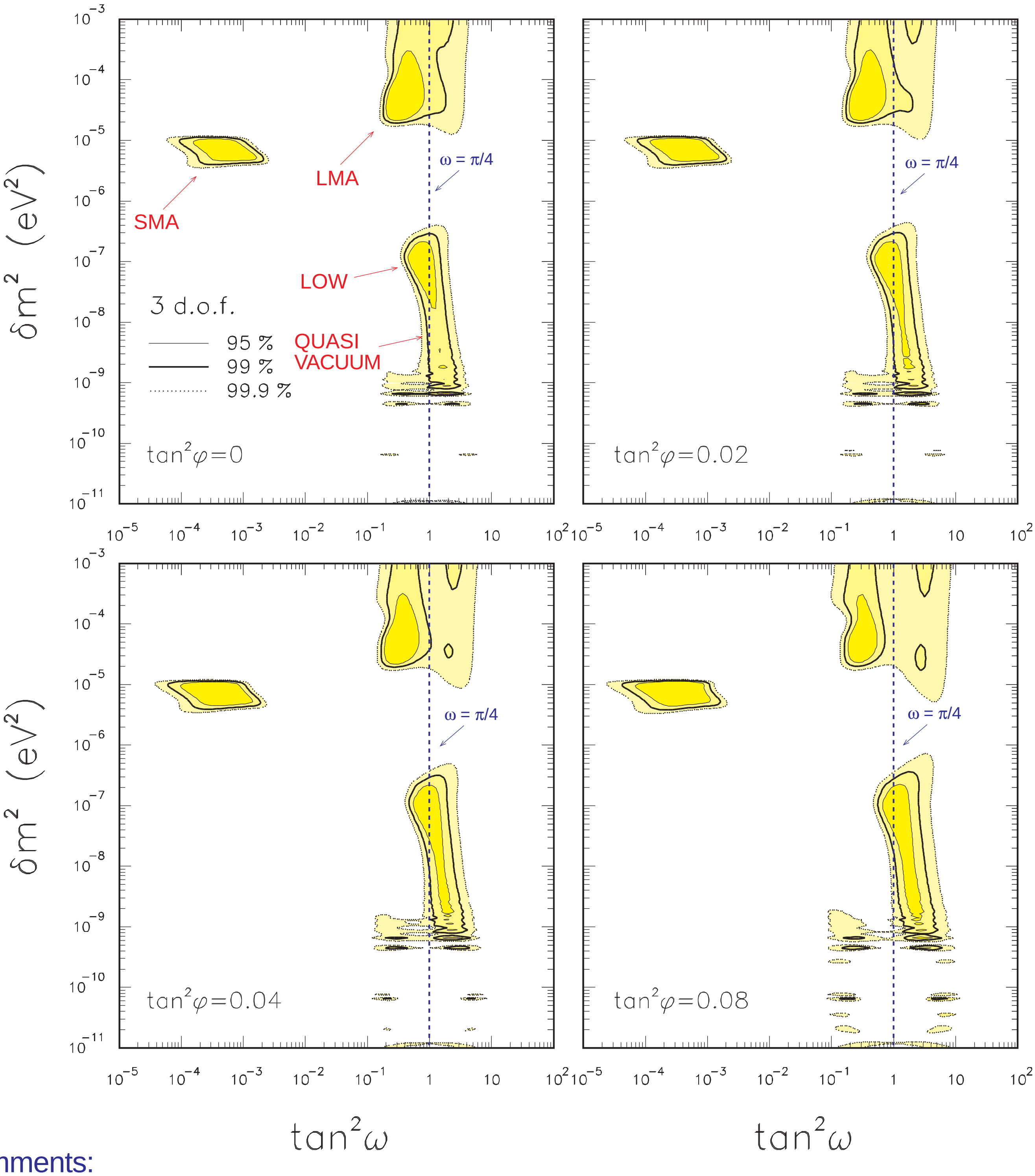
$\Downarrow$

δm^2 must be smaller than ~ 10^{-3} eV ^{2}

CHOOZ puts upper limits on $\begin{cases} \sin^2\varphi & (\leq \text{few } \%) \\ \delta m^2 & (\leq 10^{-3} \text{ eV}^2) \end{cases}$

3ν solar solutions for finite m^2 ($m^2 = 1.5 \times 10^{-3}$ eV ^{2})

total rates with constraints from Sp(D)+Sp(N)

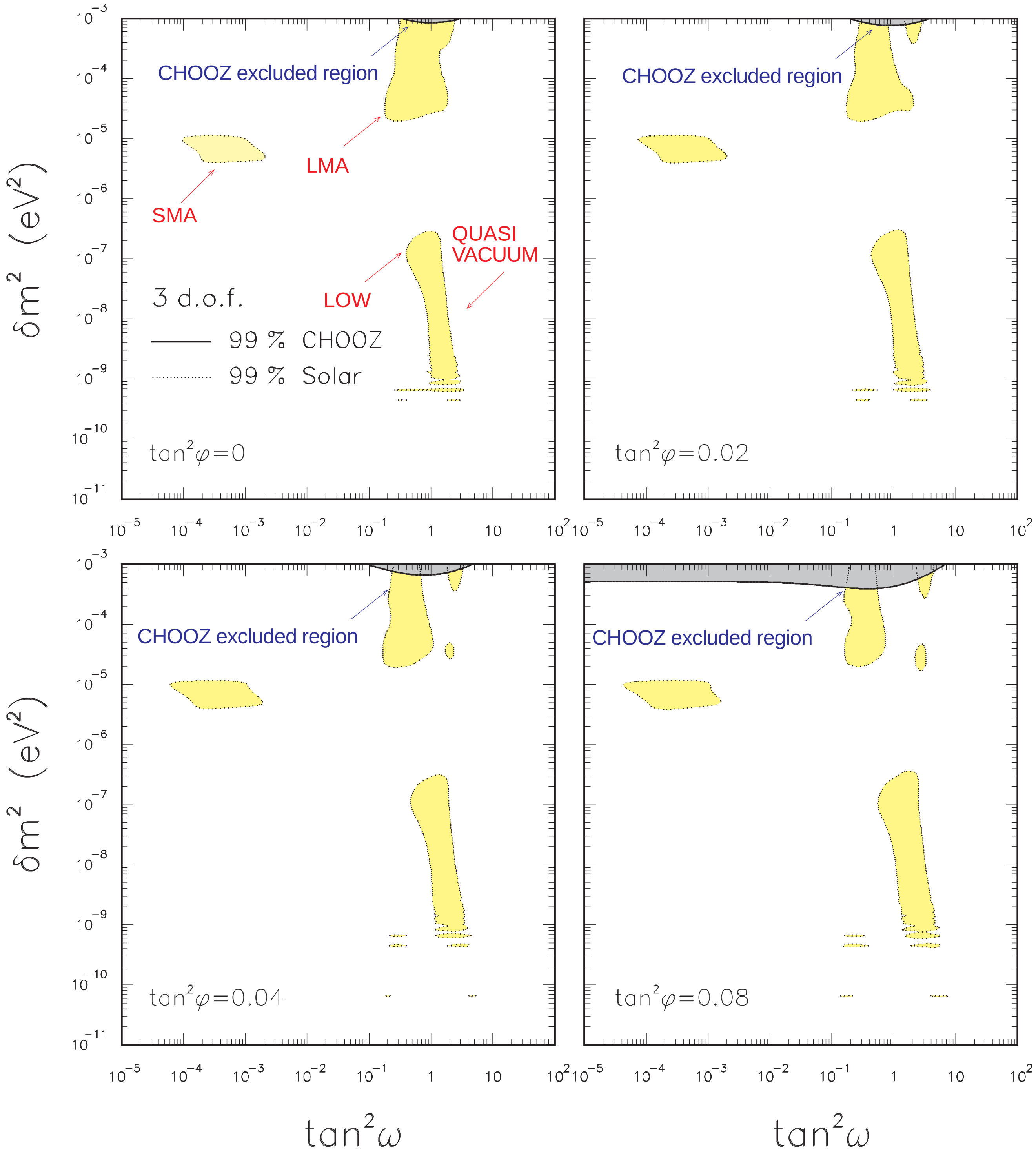


Comments:

- Removal of the $m^2 = \infty$ approximation in the solar ν analysis would not practically change the solar ν solutions
- In particular, no upper bound on δm^2 at 99% C.L. from solar ν data alone

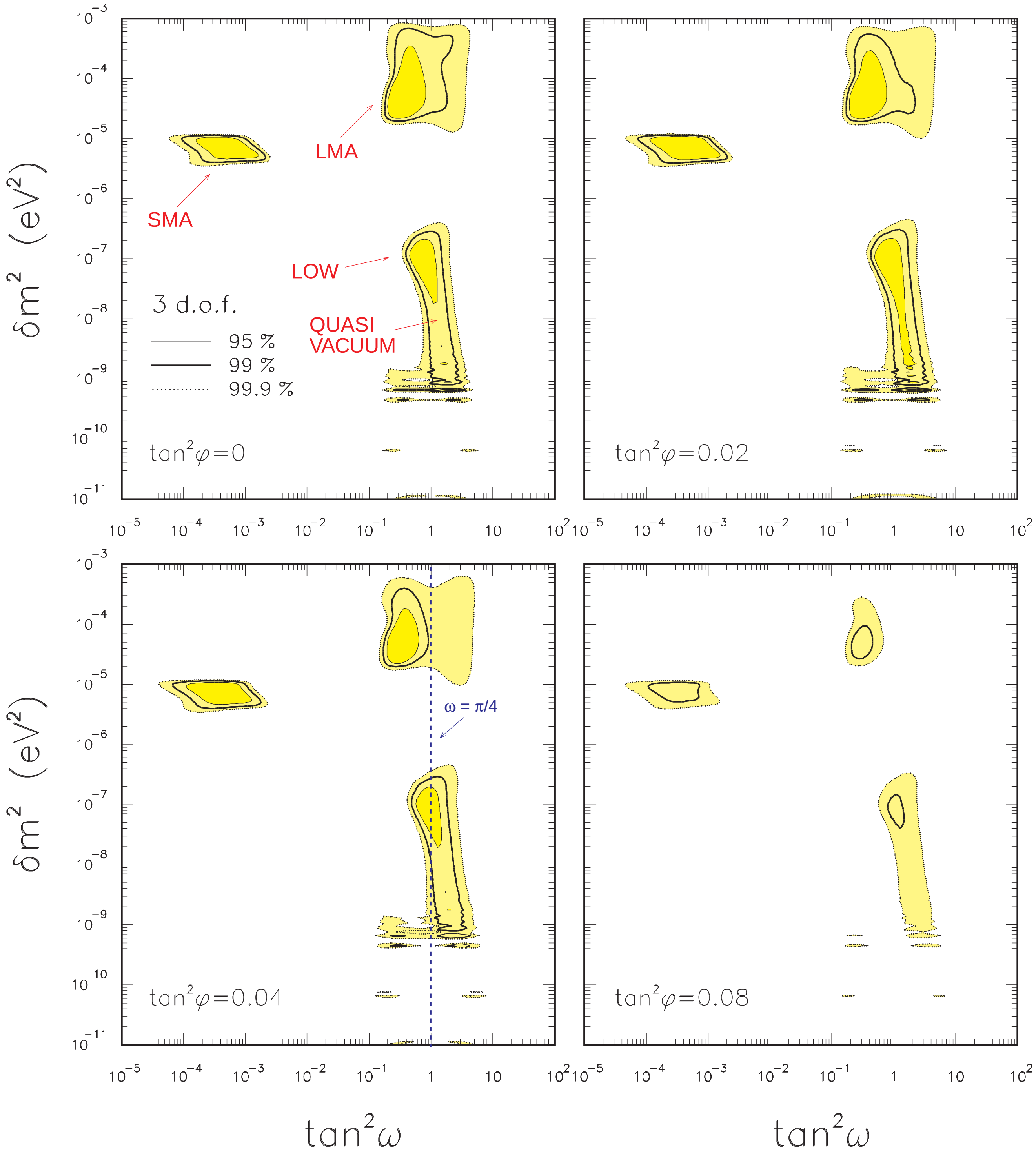
CHOOZ excluded region compared with the 3ν solar solutions

($m^2 = 1.5 \times 10^{-3}\,eV^2$)



solar + CHOOZ data

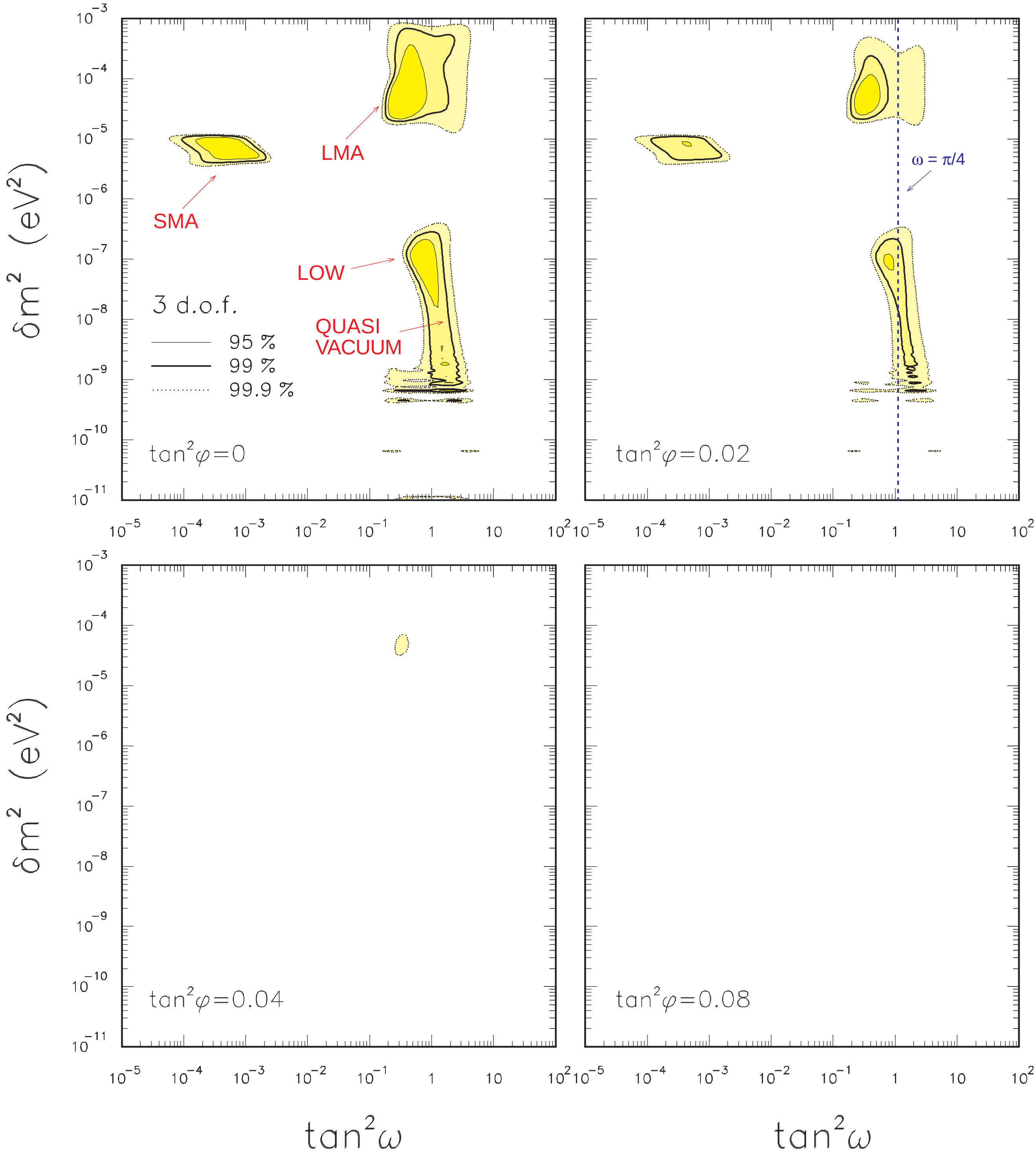
$(m^2 = 1.5 \times 10^{-3}\,\mathrm{eV}^2)$



$\Rightarrow$ Strong constraints on both δm^2 and $\tan^2\varphi$ when solar + CHOOZ data are taken into account!

solar + CHOOZ data

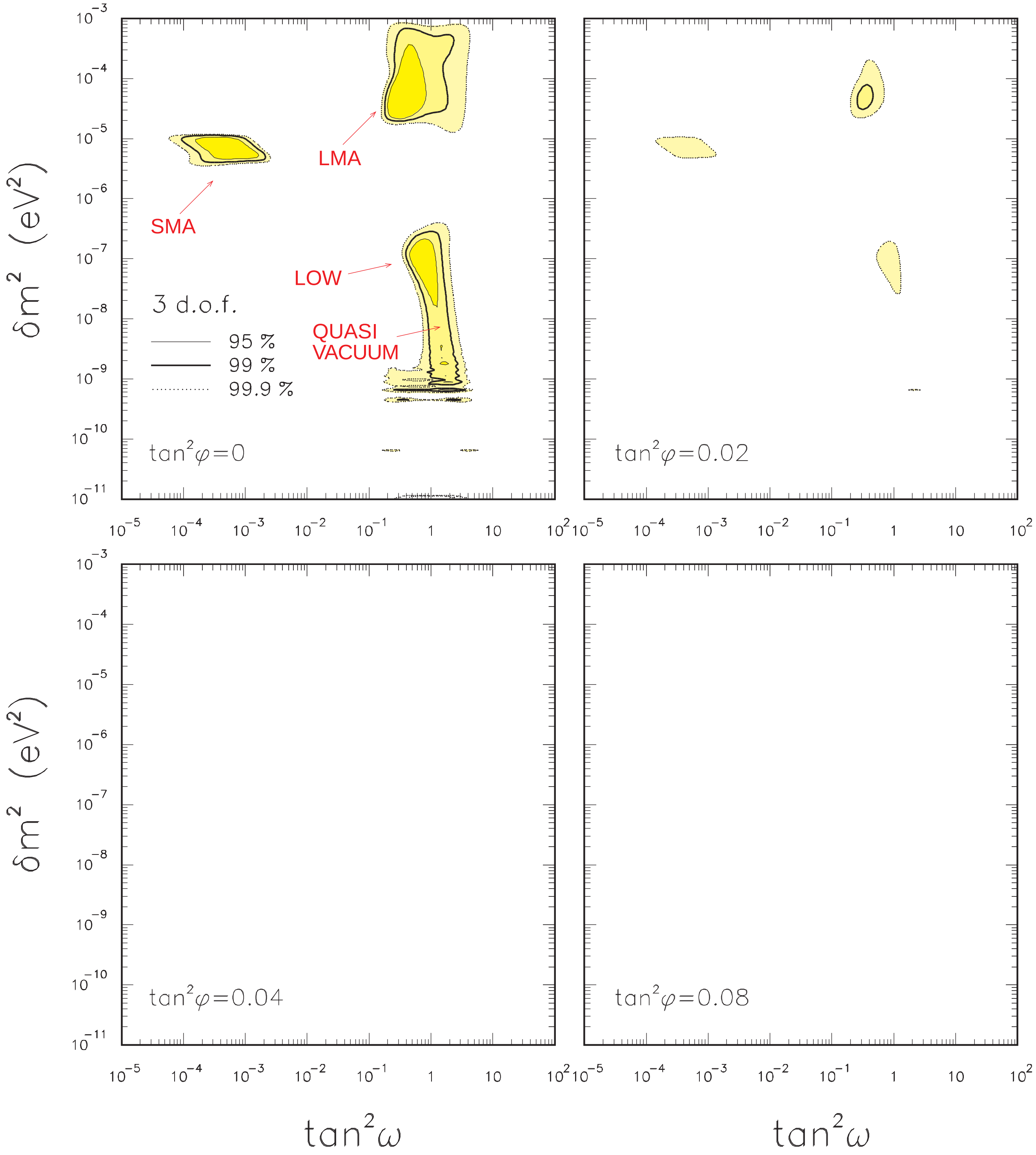
$(m^2 = 3.0 \times 10^{-3}\,\mathrm{eV}^2)$



$\Rightarrow$ More stringent constraints on $\tan^2\varphi$ for increasing m^2

solar + CHOOZ data

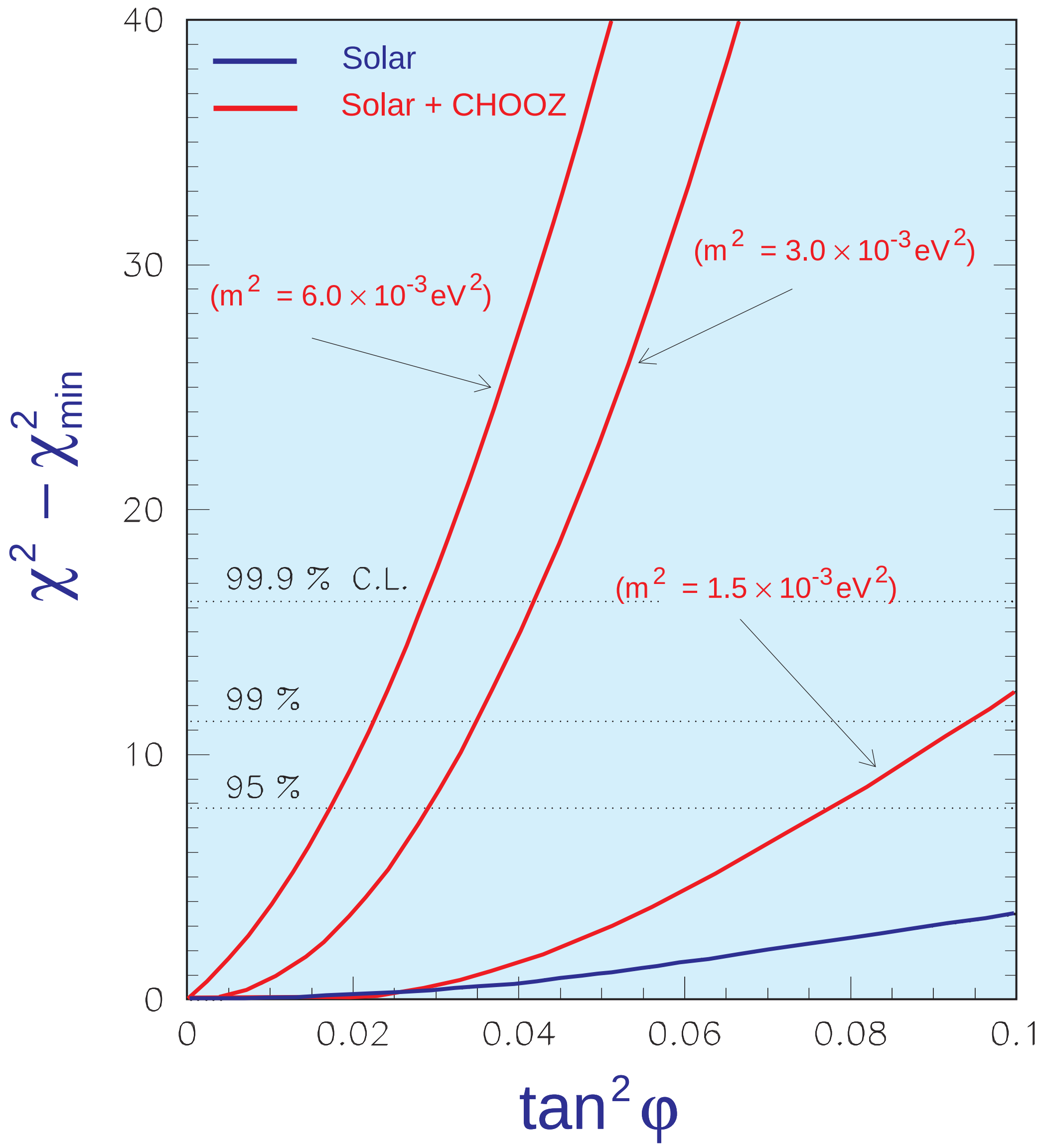
(m^2 = 6.0×10^{-3} eV ^{2})



⇒ The solar solutions in the limit $\varphi = 0$ are independent of m^2 (they correspond to a pure 2ν case)

solar + CHOOZ data

χ^2 behaviour for different values of m^2

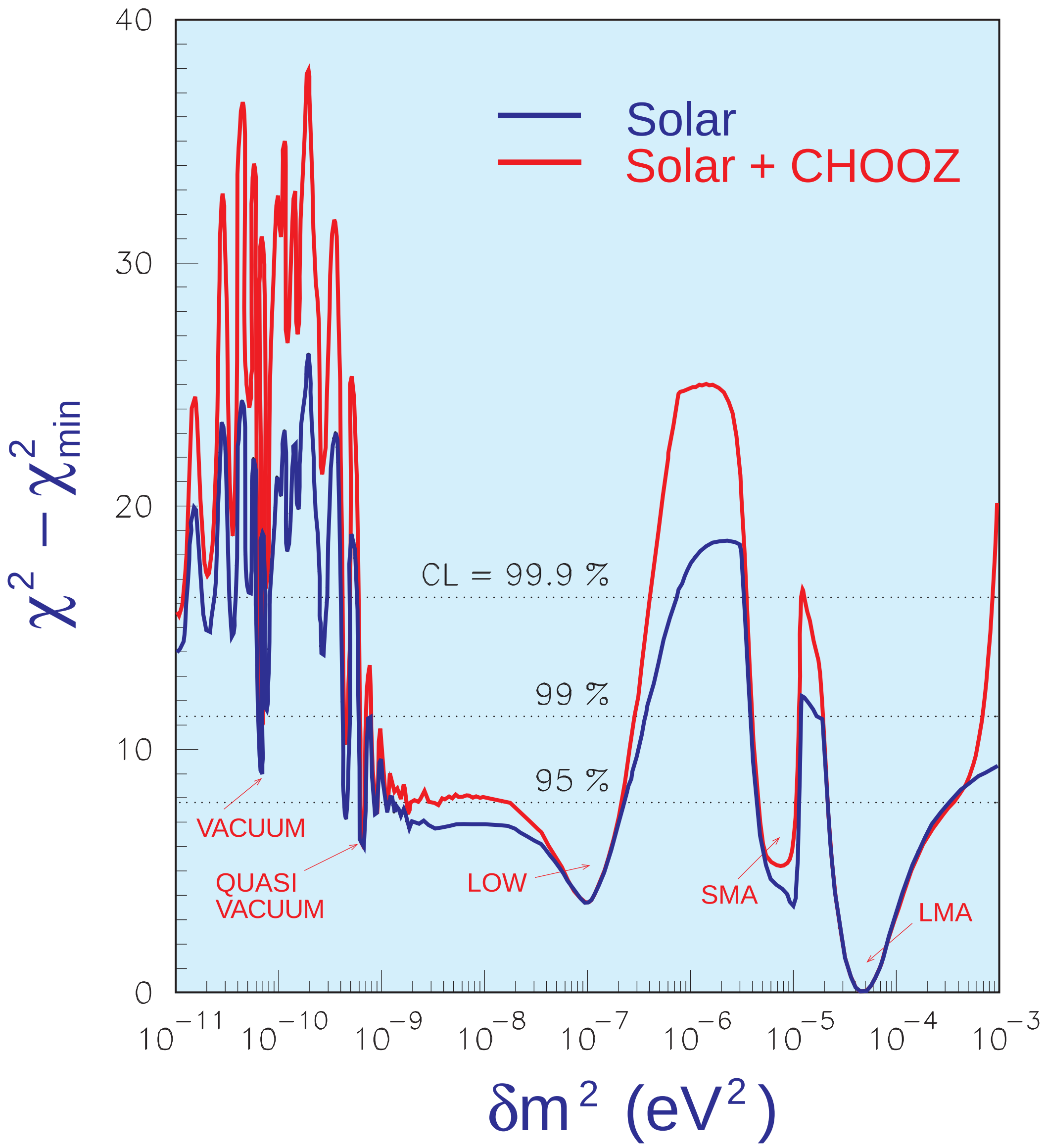


- Strong constraints on $\tan^2\varphi$ when solar + CHOOZ data are taken into account

- More stringent constraints on $\tan^2\varphi$ for increasing m^2

solar + CHOOZ data

χ^2 behaviour in terms of of δm^2 ($\tan^2\varphi$ unconstrained)



- More stringent constraints on δm^2 from solar + CHOOZ data

Summary on solar ν (in 6 points)

1. Still multiplicity of solar neutrino solutions:

   MSW: LMA preferred, LOW promising, SMA less favored
   VAC: serious problems for VAC, but interest for quasi-VAC (in the range $10^{-8} \div 10^{-9}$ eV^2)

2. The LOW solution is acceptable and can provide

   maximal $\nu_1 \leftrightarrow \nu_2$ mixing ($\omega = \pi/4$) for nonzero φ

3. Bounds on φ from solar ν data alone are loose (even more than for atmospheric data). However, their preference for

   $$s^2_\varphi \sim 0$$

   is a good sign of consistency with CHOOZ.

4. CHOOZ provides upper bounds on s^2_φ and δm^2

5. Unambiguous selection of ONE solution will not be possible in SK for several years. Much higher statistics needed (e.g., in the spectrum).

6. We absolutely need more solar ν data ! Eagerly waiting for SNO, GNO, Borexino results ..., as well as for direct tests of LMA with KAMLAND, to improve discrimination of solar ν solutions.

3ν summary

Within the “standard” 3ν interpretation of neutrino oscillation data, we have a lot of experiments or projects constraining the neutrino parameters:

masses		mixing	
δm^2	U^2_{e1}	U^2_{e2}	U^2_{e3}
	$U^2_{\mu 1}$	$U^2_{\mu 2}$	$U^2_{\mu 3}$
m^2	$U^2_{\tau 1}$	$U^2_{\tau 2}$	$U^2_{\tau 3}$

probed by solar neutrino experiments

probed by "terrestrial" neutrino experiments

- At present, most stable constraints on m^2 and $U^2_{\mu 3}$, $U^2_{\tau 3}$.
- CHOOZ tells us that U^2_{e3} must be small, and atmospheric and solar data also prefer small φ, but there is no reason for it to be zero !
- Constraints on δm^2, U^2_{e1}, U^2_{e2} depend on which solar solution is picked up.
- CHOOZ + solar data $\longrightarrow$ upper bound on δm^2
- Most important tasks for the next years:

$\Longrightarrow$ measure or constrain further U^2_{e3} (reactors & LBL)

$\Longrightarrow$ reduce the multiplicity of solar ν solutions

$\Longrightarrow$ check and (dis)prove non-standard interpretations